Tim Linden

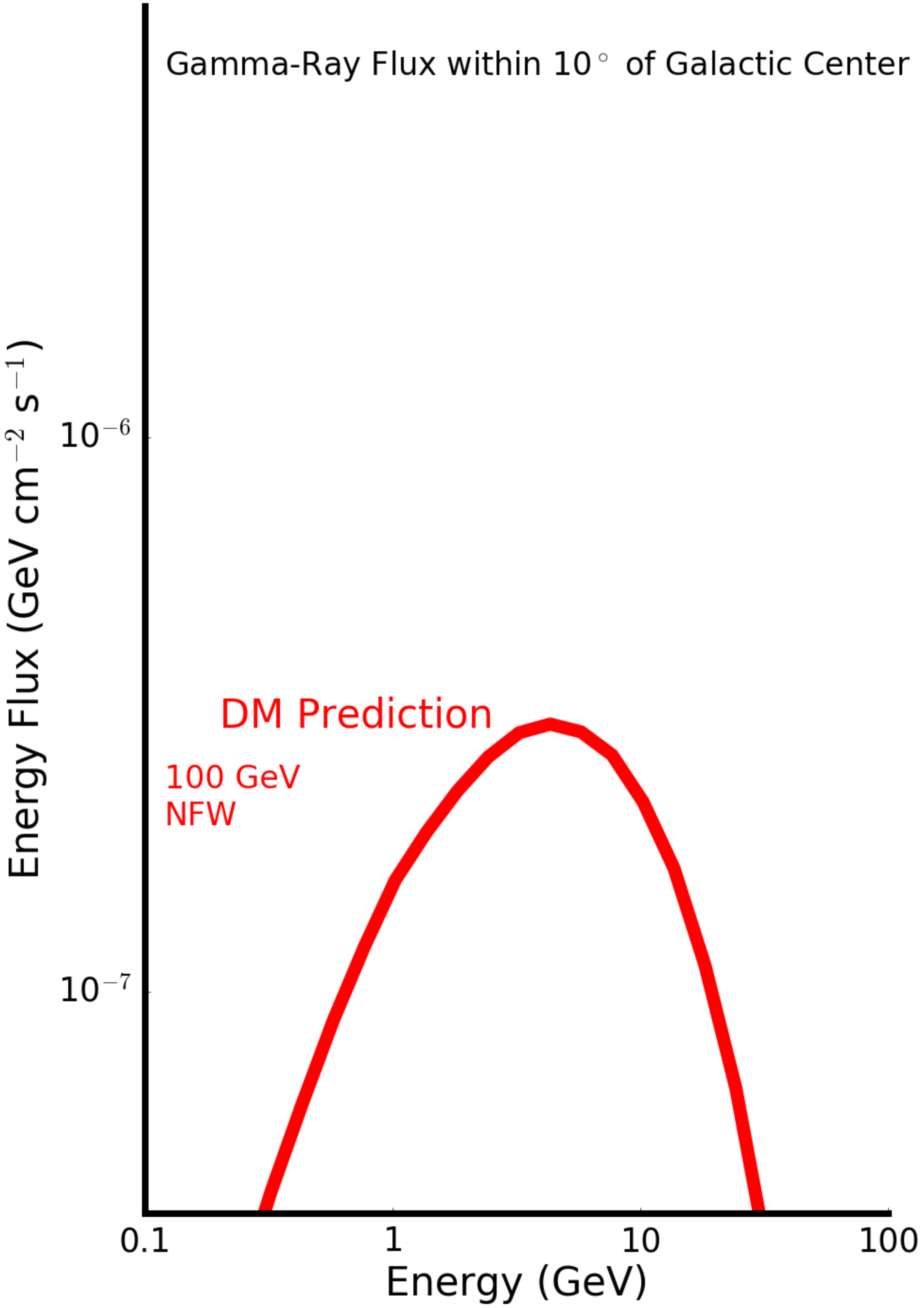
Antihelium from Dark Matter Annihilation

Stockholms universitet

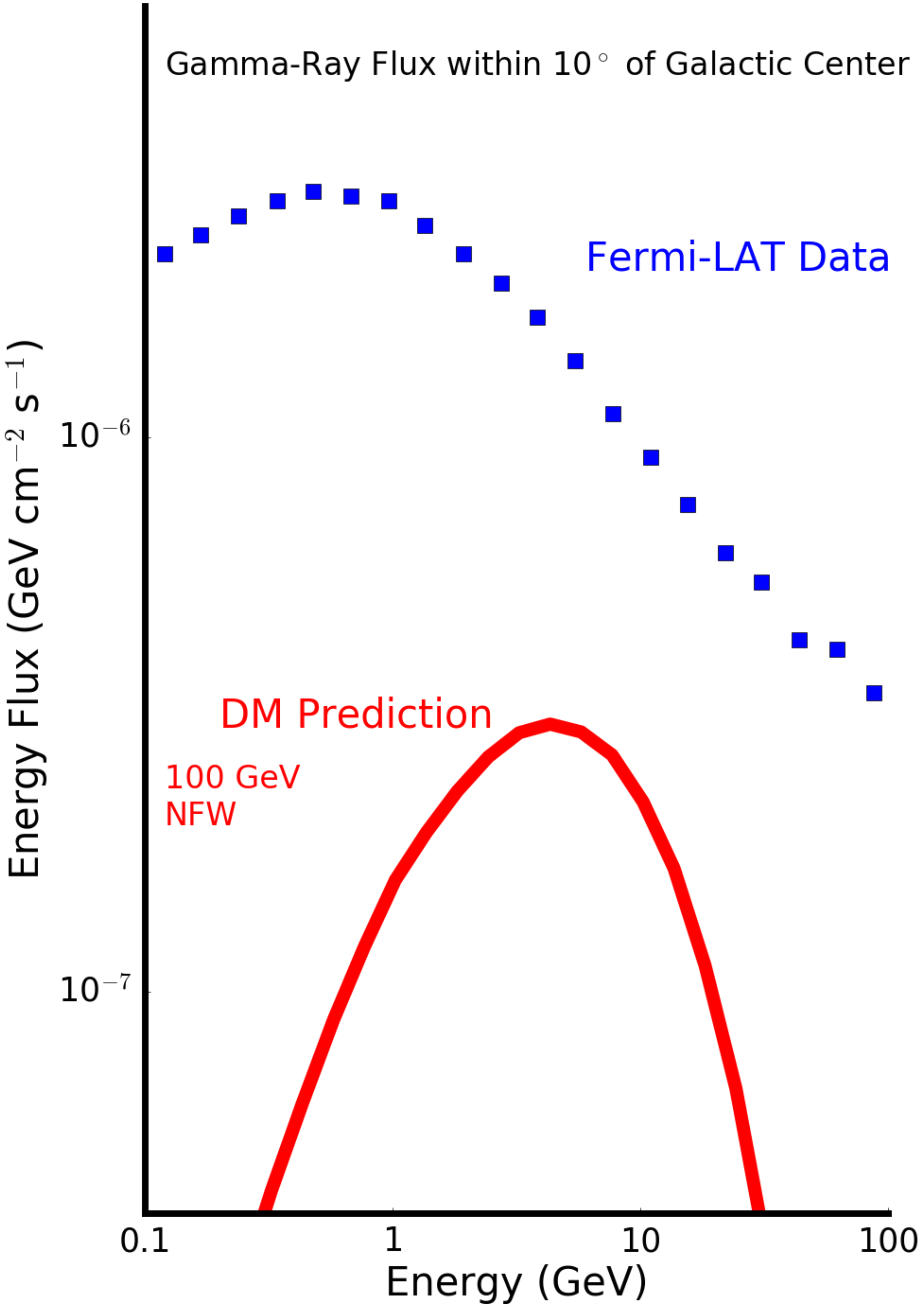

Photon
Positrons
+
-
Electrons
Leptons
Neutrinos
Quarks
-
Antiprotons
+
Protons
WIMP
Bosons

Thermal Cross-Section
1 Observed Photon Within 10° of Galactic Center
$<\sigma v>$ (cm^{3}s^{-1})
10^{-25}
10^{-26}
10^{-27}
10^{-28}
10^{-29}
10^{-30}
10^{-31}
10^{-32}
10^{-33}
10
100
1000
10^4
10^5
Dark Matter Mass (GeV)

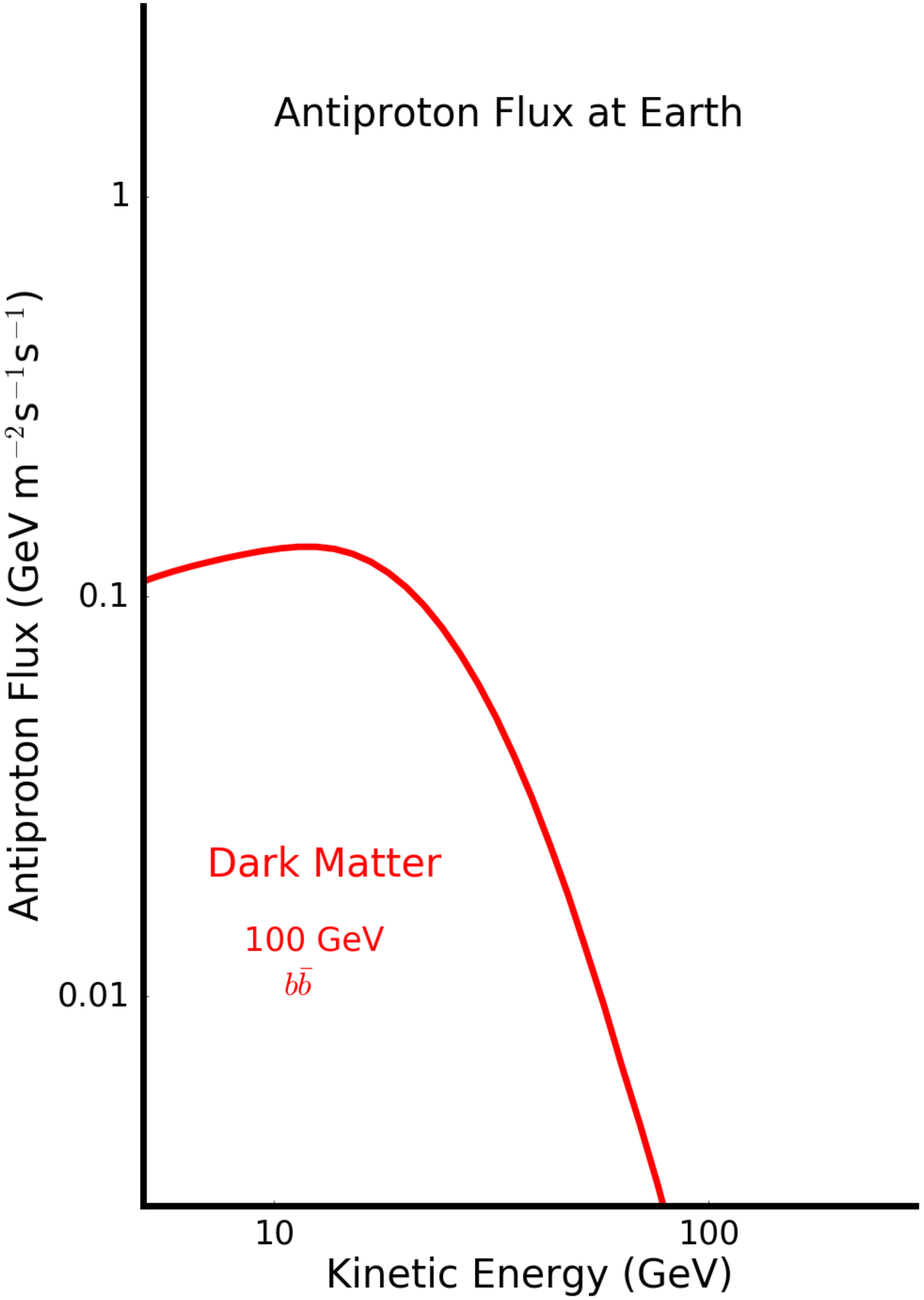
Thermal WIMPs and the Story of Tantalus



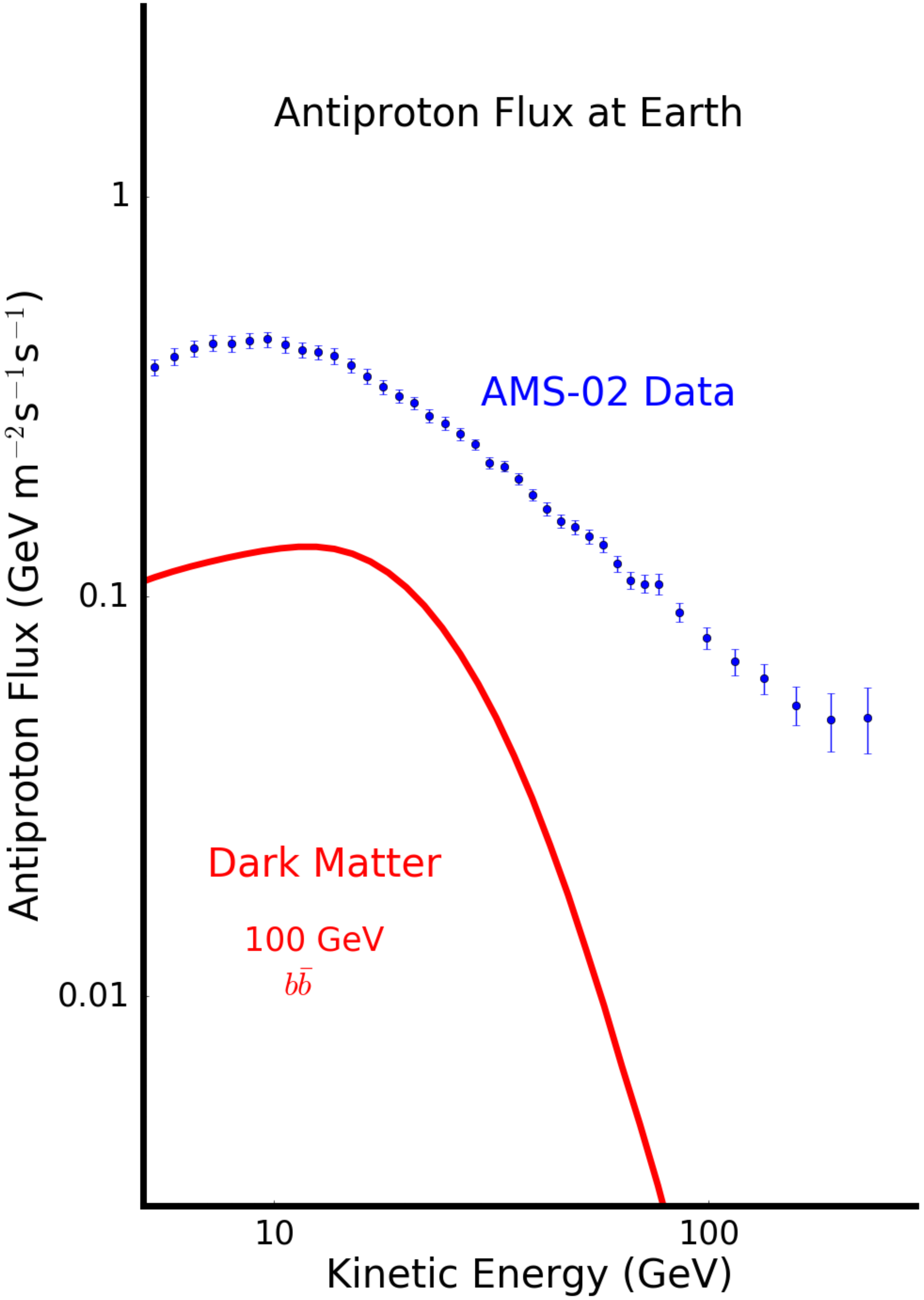
Thermal WIMPs and the Story of Tantalus



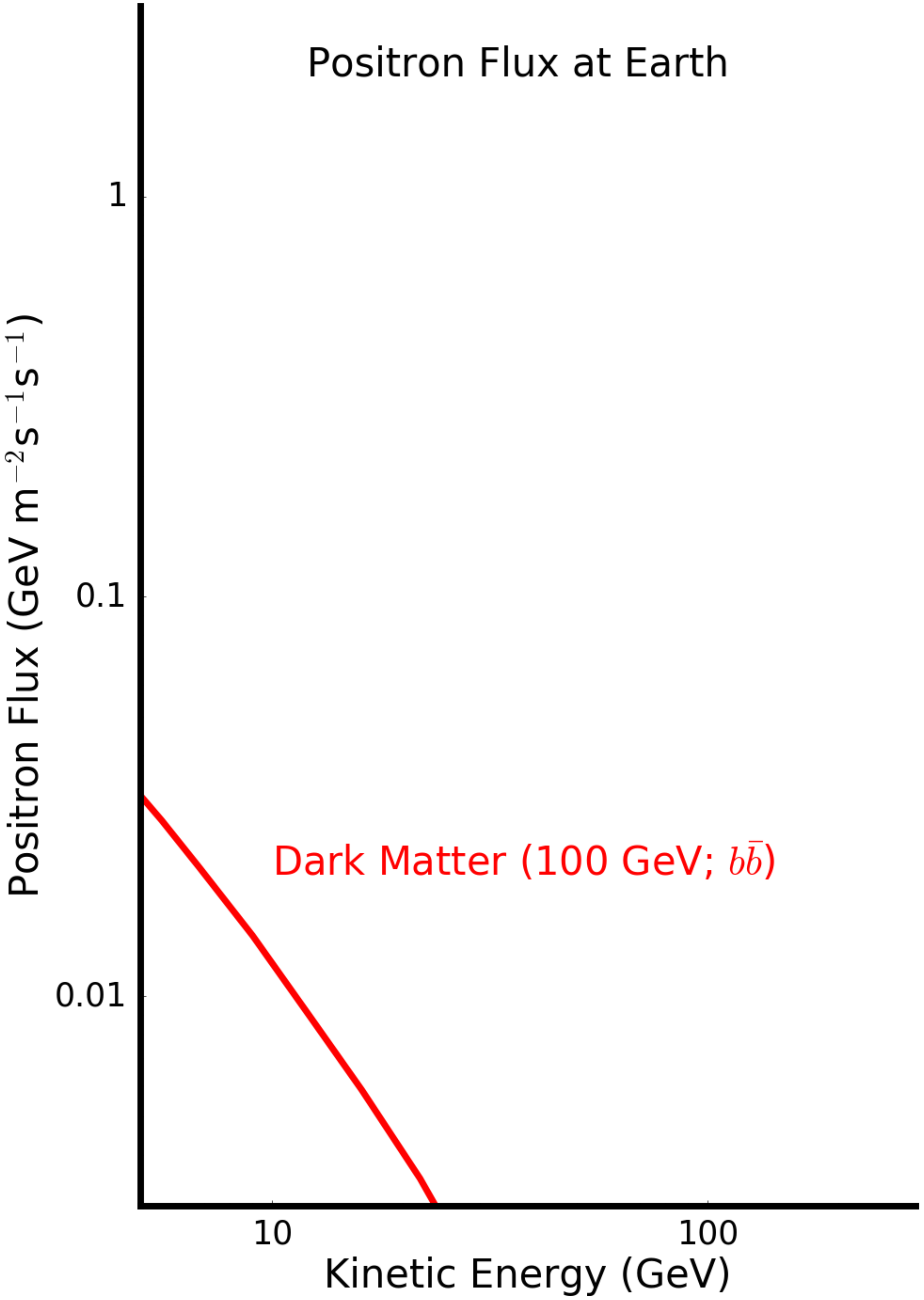
Thermal WIMPs and the Story of Tantalus



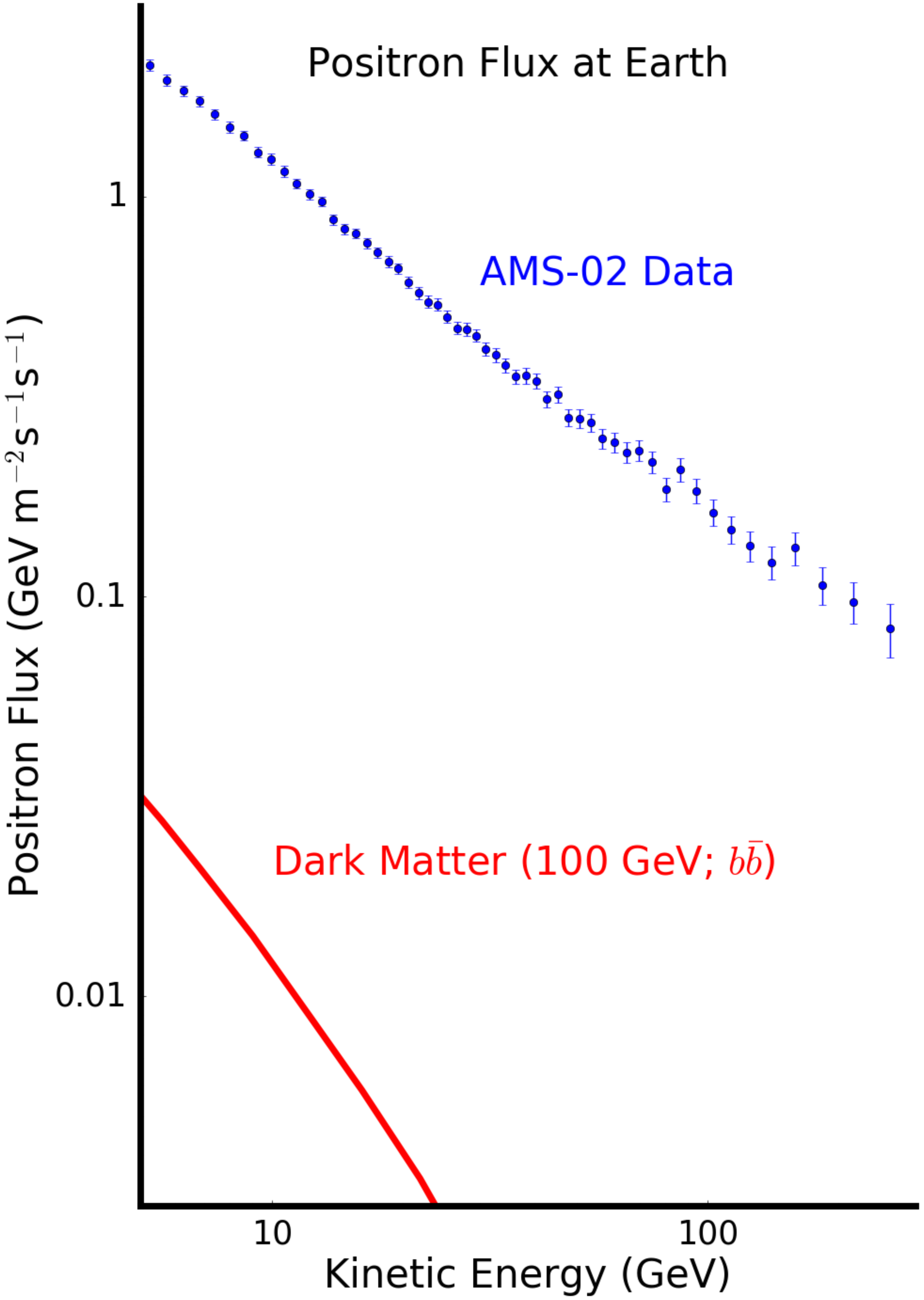
Thermal WIMPs and the Story of Tantalus



Thermal WIMPs and the Story of Tantalus



Thermal WIMPs and the Story of Tantalus




Small Dark Matter Signal
Small Astrophysical Background
Large Dark Matter Signal
Small Astrophysical Background
Small Dark Matter Signal
Large Astrophysical Background
Large Dark Matter Signal
Large Astrophysical Background
Specificity (DM Flux / Astrophysics Flux)
Fraction of Dark Matter Flux


Small Dark Matter Signal
Small Astrophysical Background
Large Dark Matter Signal
Small Astrophysical Background
Easy
Small Dark Matter Signal
Large Astrophysical Background
Large Dark Matter Signal
Large Astrophysical Background
Specificity (DM Flux / Astrophysics Flux)
Fraction of Dark Matter Flux


Small Dark Matter Signal
Small Astrophysical Background
Large Dark Matter Signal
Small Astrophysical Background
Easy
Easy
Small Dark Matter Signal
Large Astrophysical Background
Large Dark Matter Signal
Large Astrophysical Background
Specificity (DM Flux / Astrophysics Flux)
Fraction of Dark Matter Flux

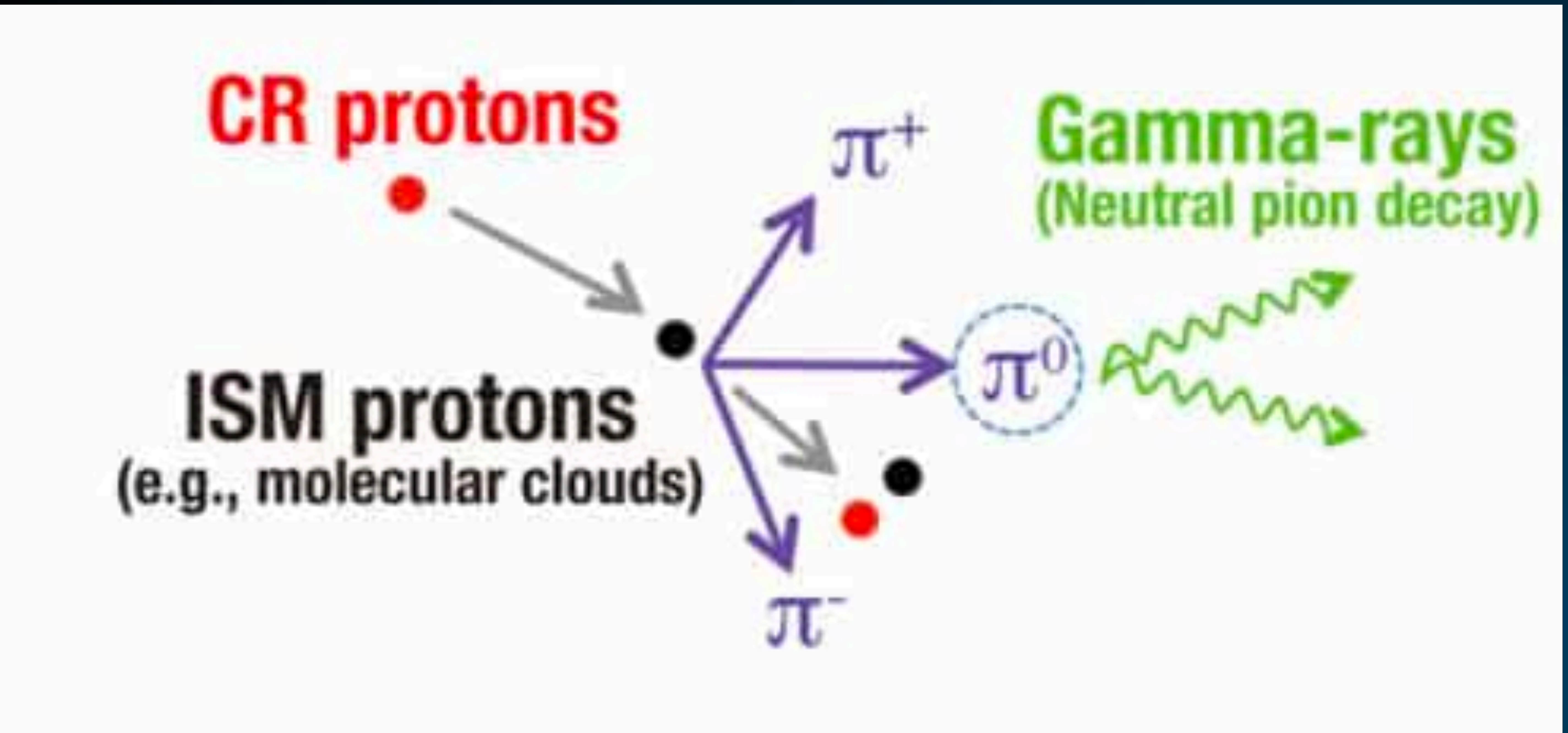

Specificity (DM Flux / Astrophysics Flux)
Small Dark Matter Signal
Small Astrophysical Background
Large Dark Matter Signal
Small Astrophysical Background
Easy
Easy
Hard
Small Dark Matter Signal
Large Astrophysical Background
Large Dark Matter Signal
Large Astrophysical Background
Fraction of Dark Matter Flux


Small Dark Matter Signal
Small Astrophysical Background
Acceptable
Large Dark Matter Signal
Small Astrophysical Background
Easy
Easy
Small Dark Matter Signal
Large Astrophysical Background
Hard
Large Dark Matter Signal
Large Astrophysical Background
Specificity (DM Flux / Astrophysics Flux)
Fraction of Dark Matter Flux


Anti-Nuclei
Antiprotons
Gamma-Rays / Positrons
Specificity (DM Flux / Astrophysics Flux)
Fraction of Dark Matter Flux

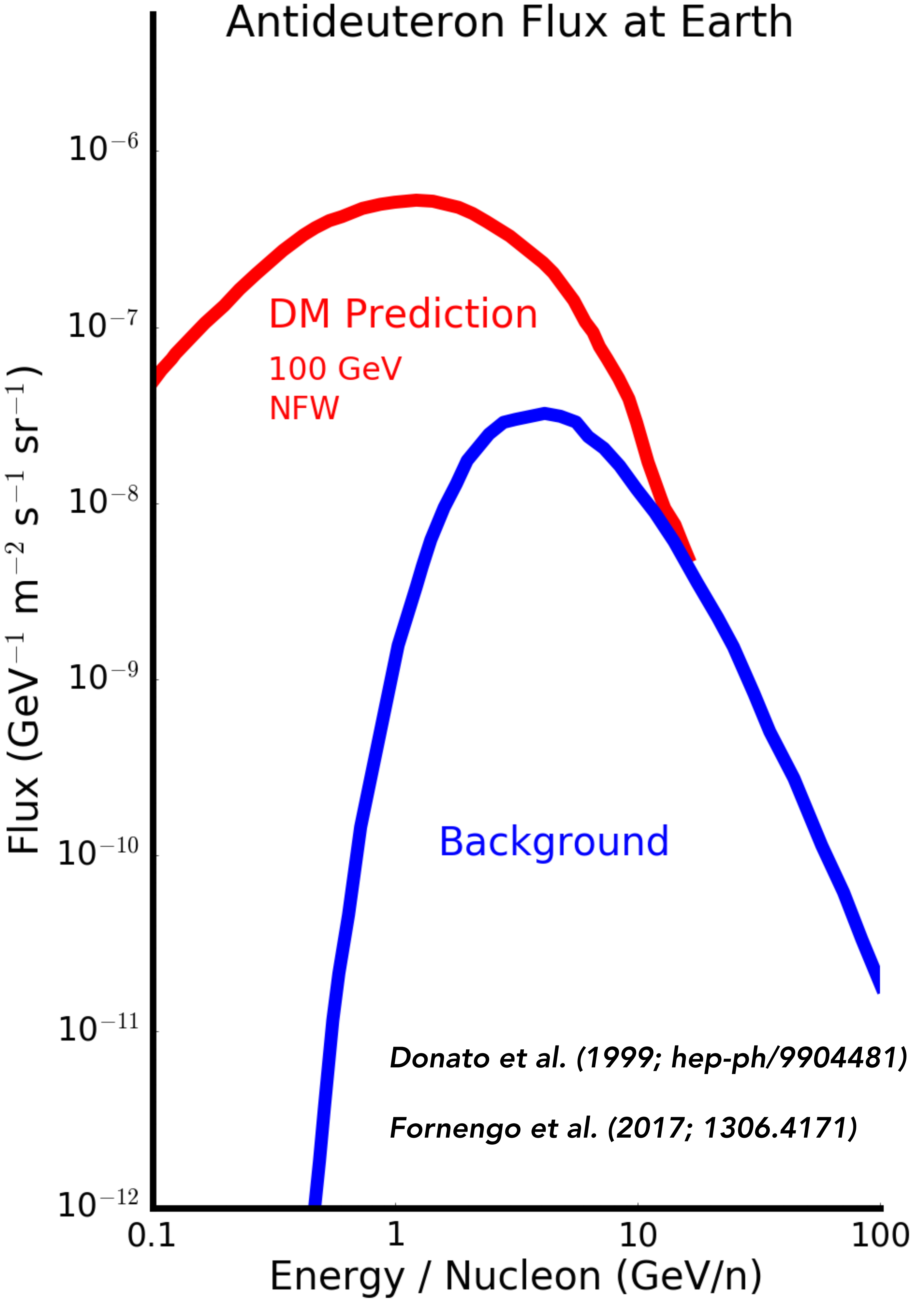

NOT MY DEPARTMENT
BRILLIANT!
NOT MY PROBLEM

AntiNuclei: A Clean Search Strategy

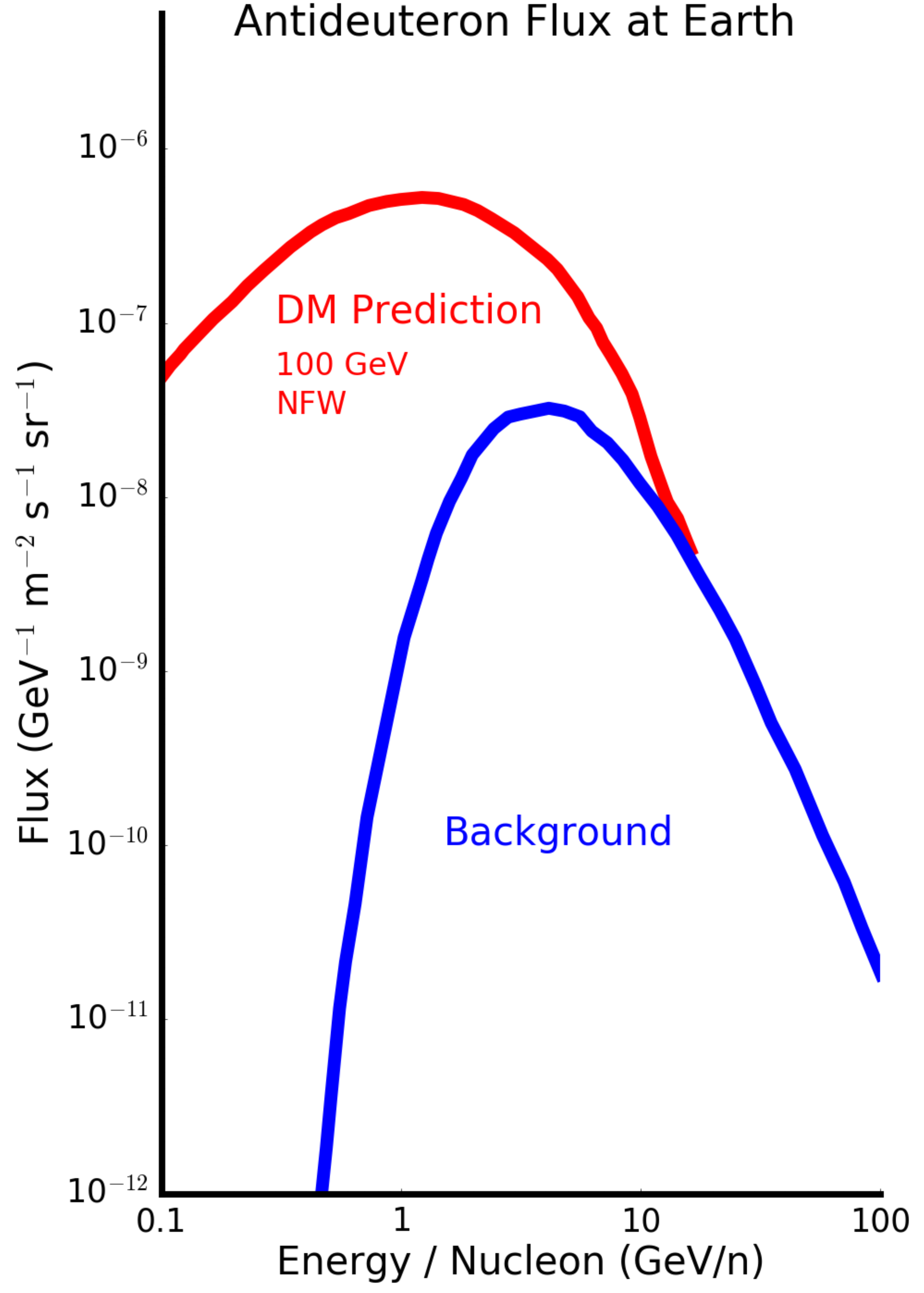


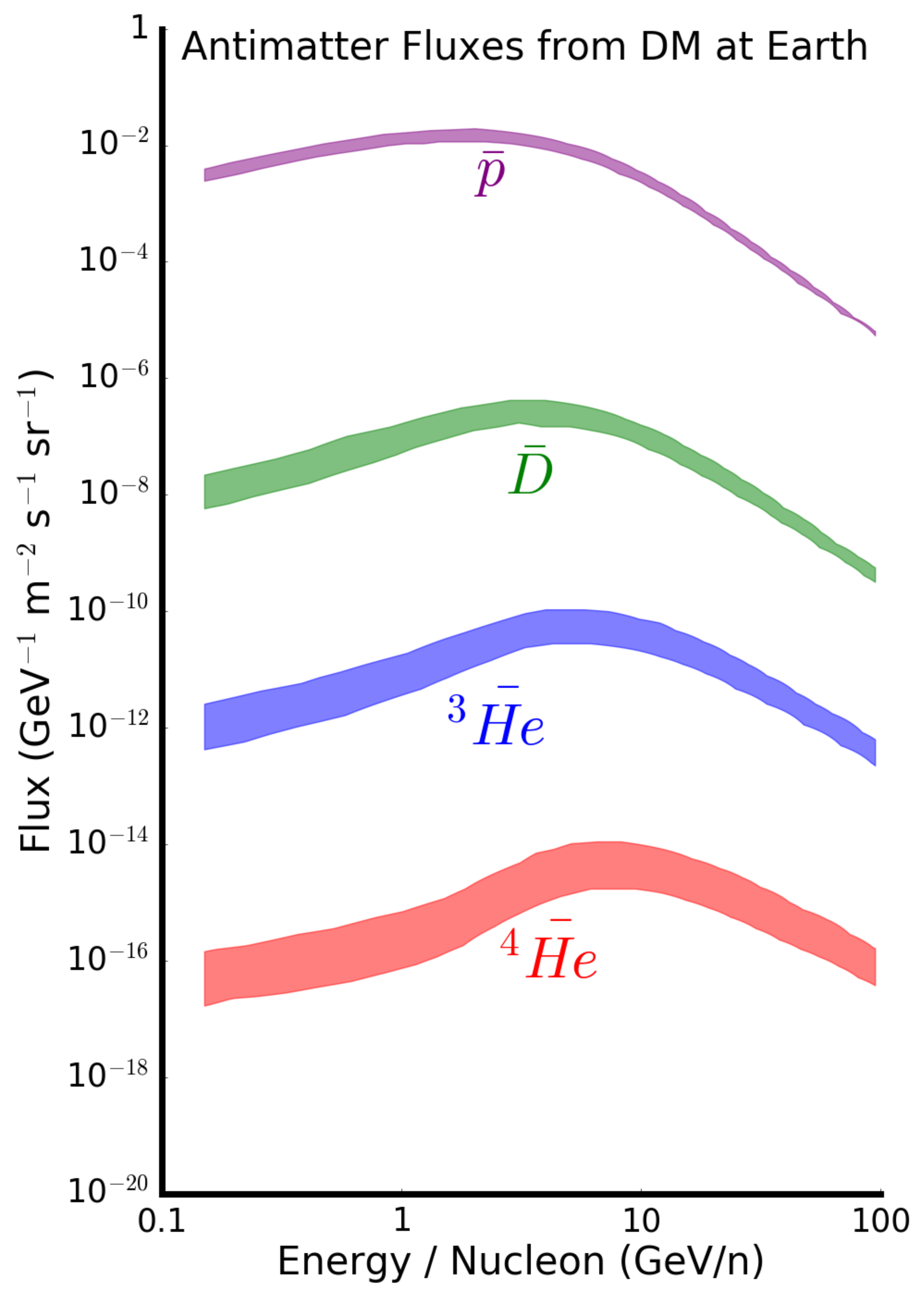
Astrophysical secondary production has a **large momentum** in the lab (galactic) frame!

Antinuclei are heavy - **inherit this momentum.**

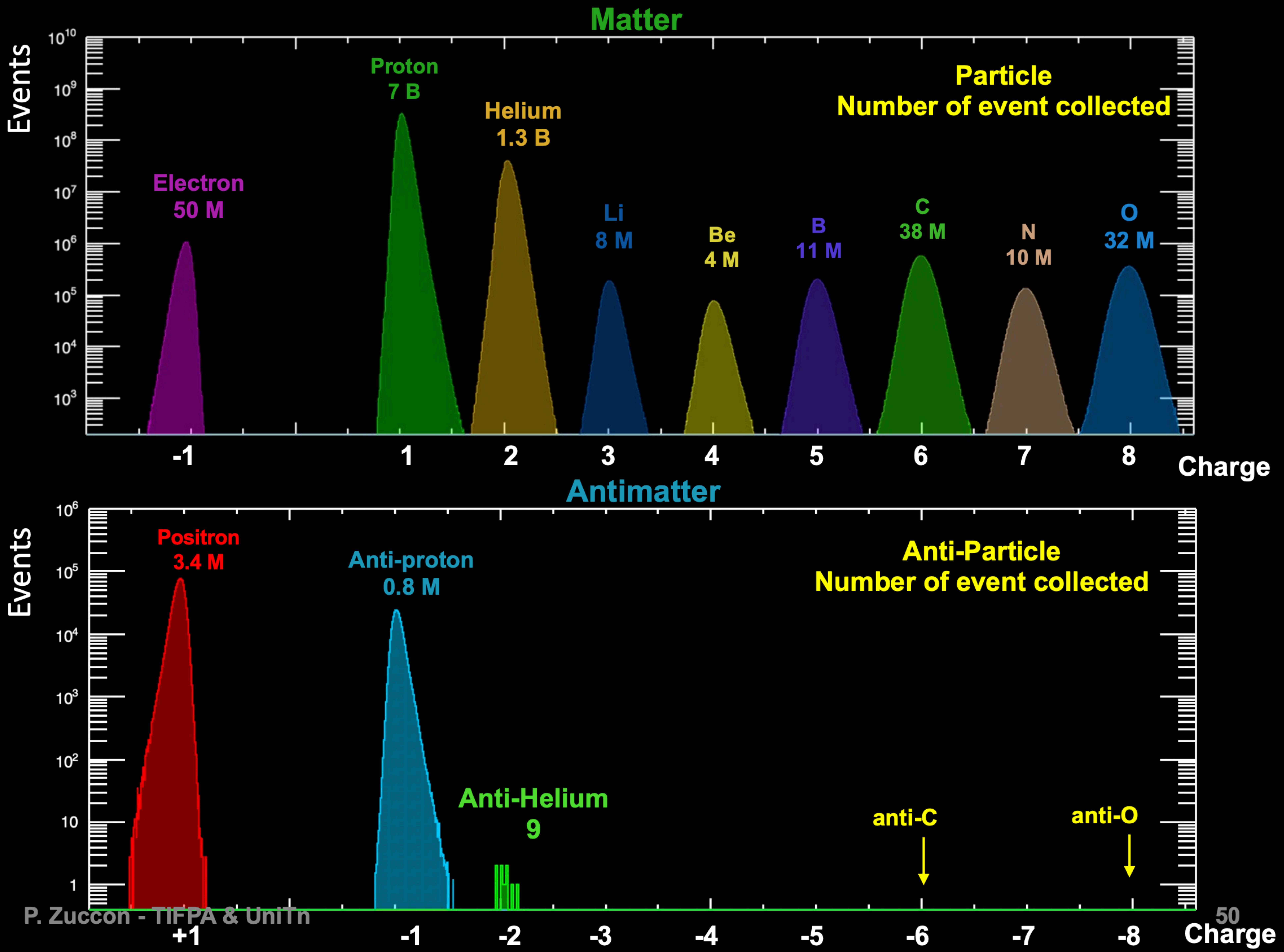


AntiNuclei: A Clean Search Strategy





Tentative Evidence for Antinuclei



Matter
Events
Particle
Number of event collected
Electron
50 M
Proton
7 B
Helium
1.3 B
Li
8 M
Be
4 M
B
11 M
C
38 M
N
10 M
O
32 M
-1
1
2
3
4
5
6
7
8
Charge
Antimatter
Events
Anti-Particle
Number of event collected
Positron
3.4 M
Anti-proton
0.8 M
Anti-Helium
9
anti-C
anti-O
+1
-1
-2
-3
-4
-5
-6
-7
-8
Charge

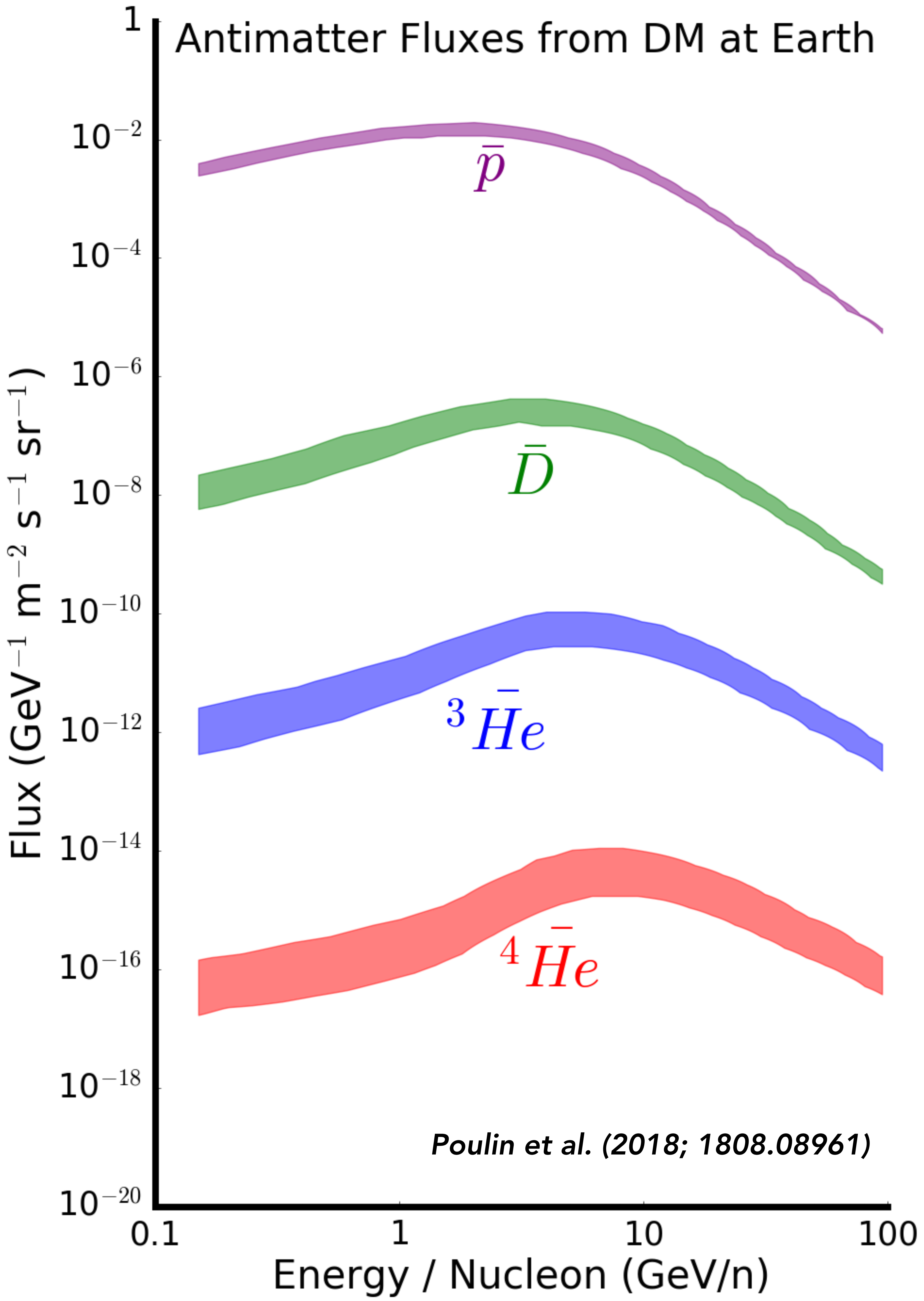
Chasing an AntiHelium Signal

1.) Coalescence Rates (1401.2461)

- Astrophysical Acceleration (2001.08749)

2.) Lambda_b Enhancement (2006.16251, 2106.00053)

3.) Strongly Coupled Dark Sectors (2211.00025)



$$E_A \frac{d^3 N_A}{dp_A^3} = B_A \left(E_{\bar{\mathrm{p}}} \frac{d^3 N_{\bar{\mathrm{p}}}}{dp_{\bar{\mathrm{p}}}^3} \right)^Z \left(E_{\bar{\mathrm{n}}} \frac{d^3 N_{\bar{\mathrm{n}}}}{dp_{\bar{\mathrm{n}}}^3} \right)^{A-Z}$$

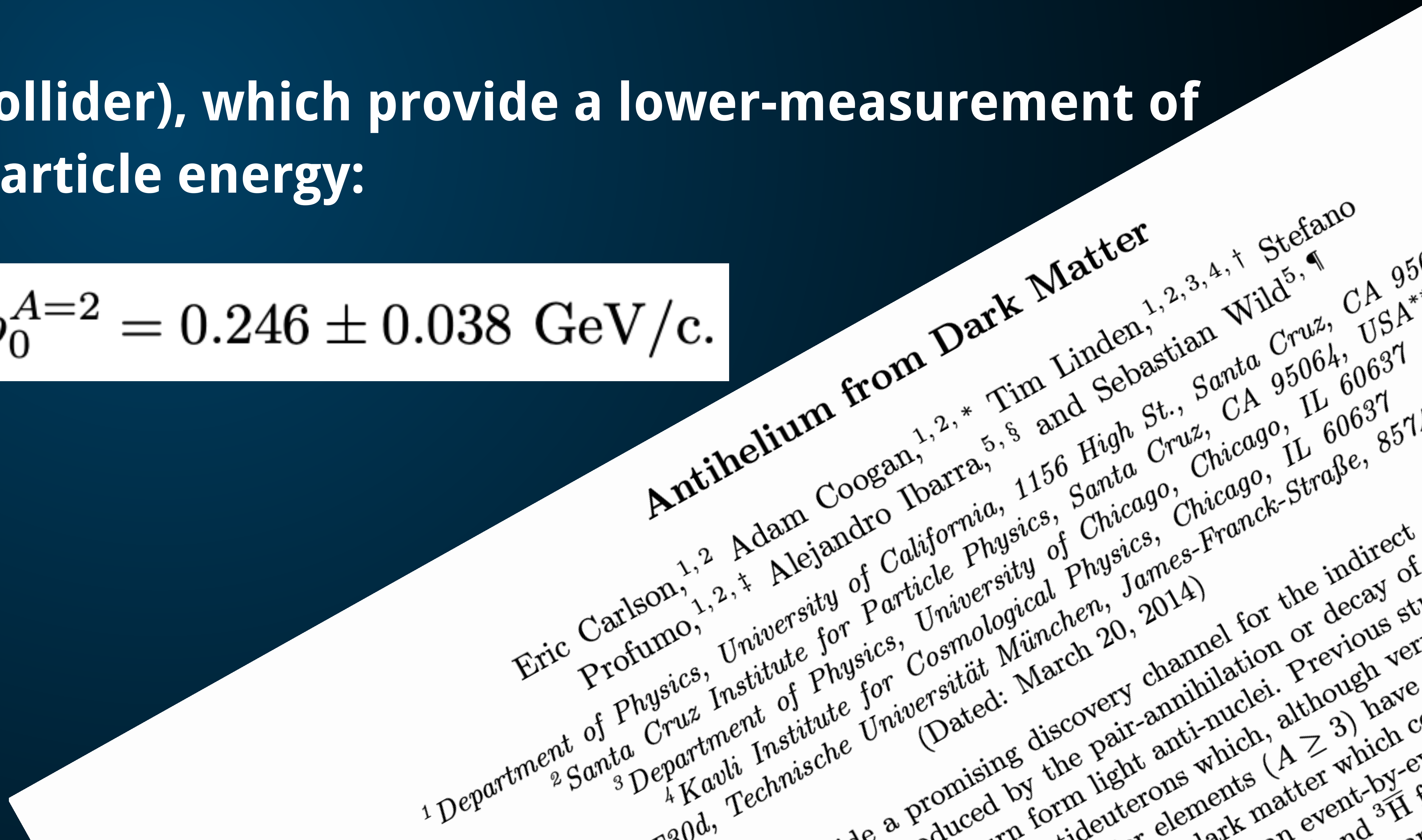
Key Insight - Coalescence Momentum for Antihelium Should Be Larger

While particle coalescence is hard to measure, the inverse process (fragmentation) is easier to measure. Helium's binding energy significantly exceeds deuteriums

$$p_0^{A=3} = \sqrt{B_{^3\overline{He}}/B_{\bar{D}}}\; p_0^{A=2} = 0.357 \pm 0.059 \text{ GeV/c}.$$

Can also use Heavy ion results (Berkeley Collider), which provide a lower-measurement of the coalescence momentum at a specific particle energy:

$$p_0^{A=3} = 1.28\; p_0^{A=2} = 0.246 \pm 0.038 \text{ GeV/c}.$$


Antihelium from Dark Matter

Eric Carlson,[1,2] Adam Coogan,[1,2,*] Tim Linden,[1,2,3,4,†] Stefano Profumo,[1,2,‡] Alejandro Ibarra,[5,§] and Sebastian Wild[5,¶]

[1] Department of Physics, University of California, 1156 High St., Santa Cruz, CA 95

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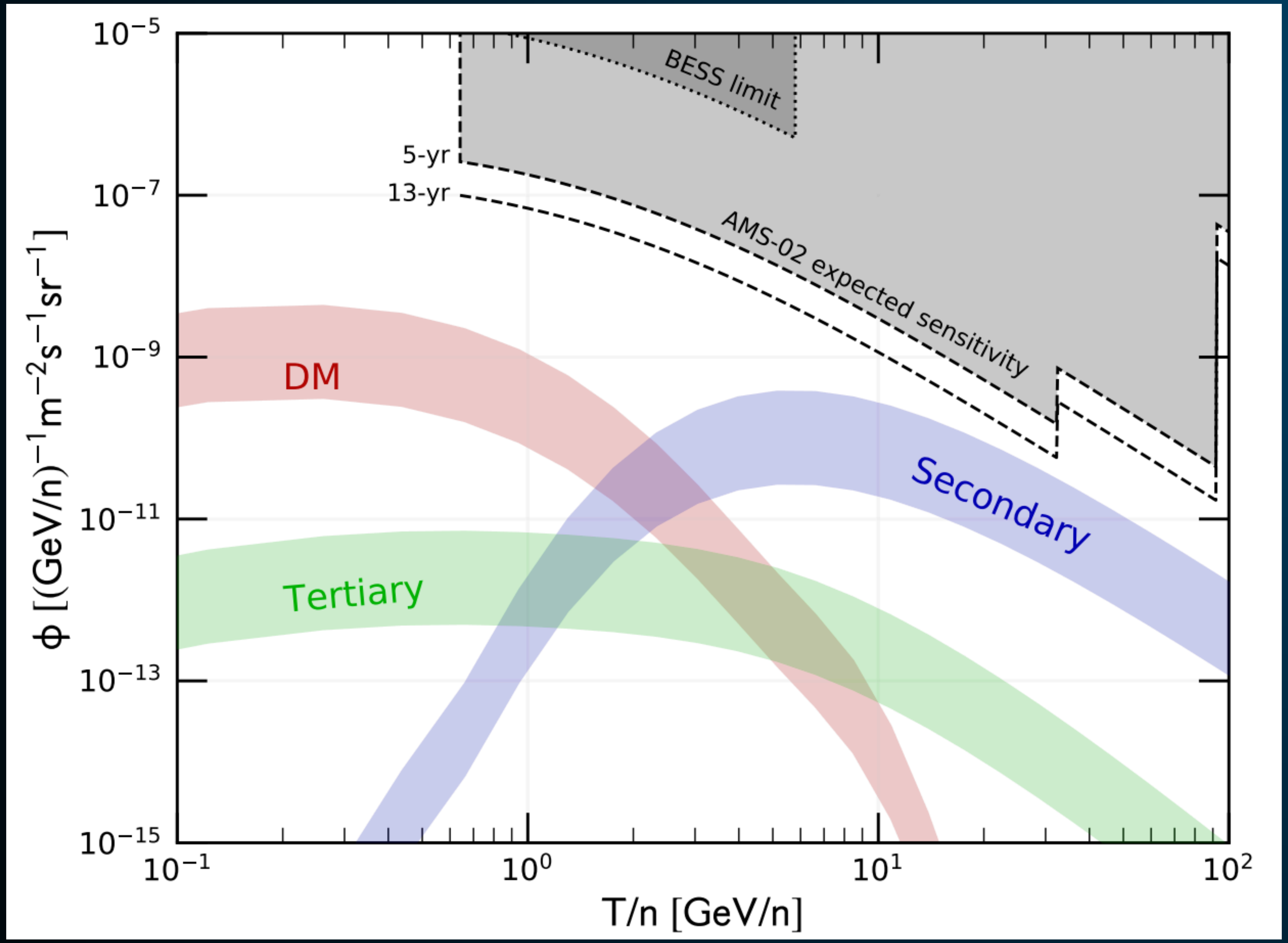
Technische Universität München, James-Franck-Straße, 857

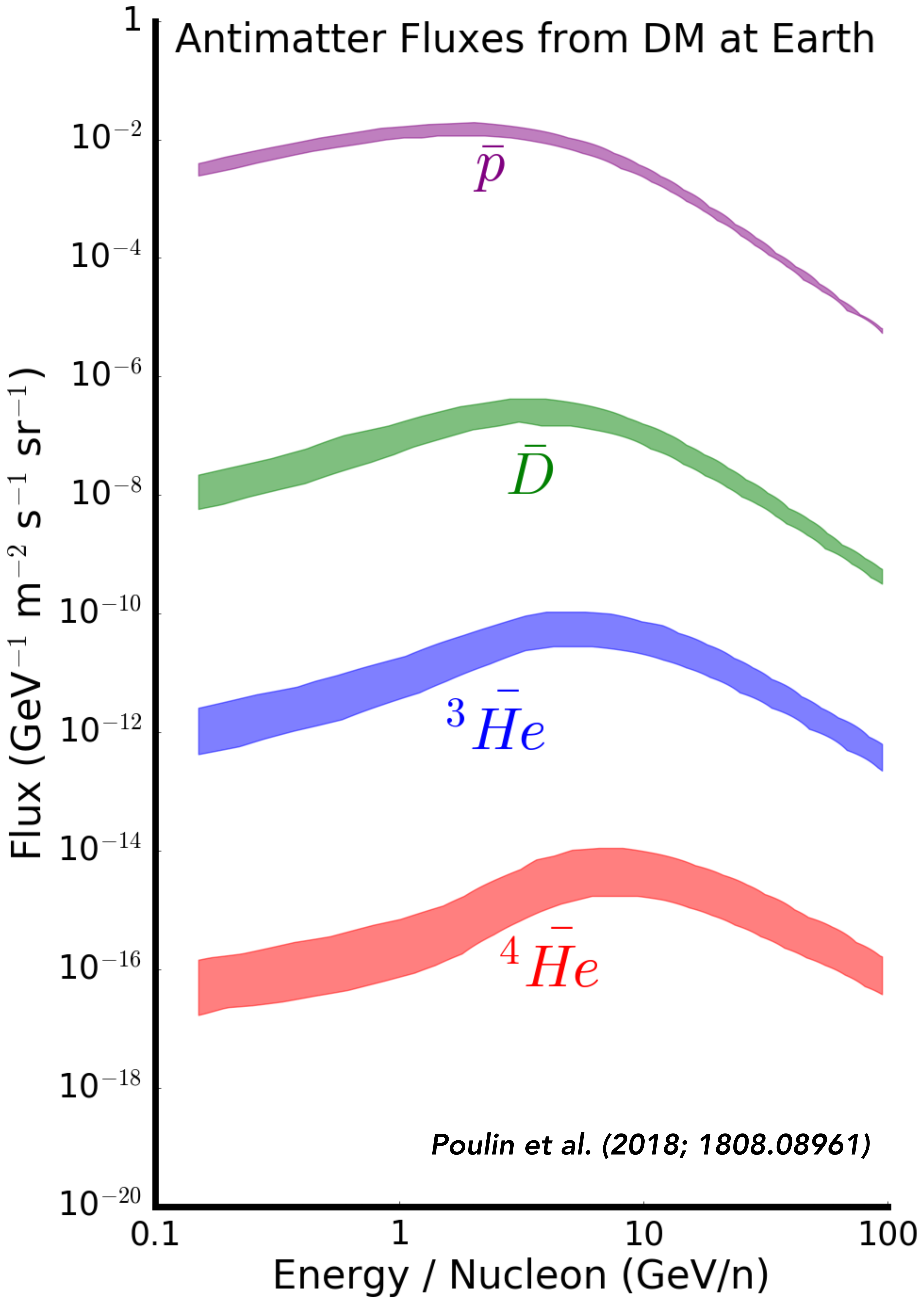
(Dated: March 20, 2014)

Coalescence Models - Expected Helium Flux

Using more realistic estimates for the anti helium coalescence momentum produces a boosted anti helium flux, especially at low energies.



Korsmeier (2017; 1711.08465)



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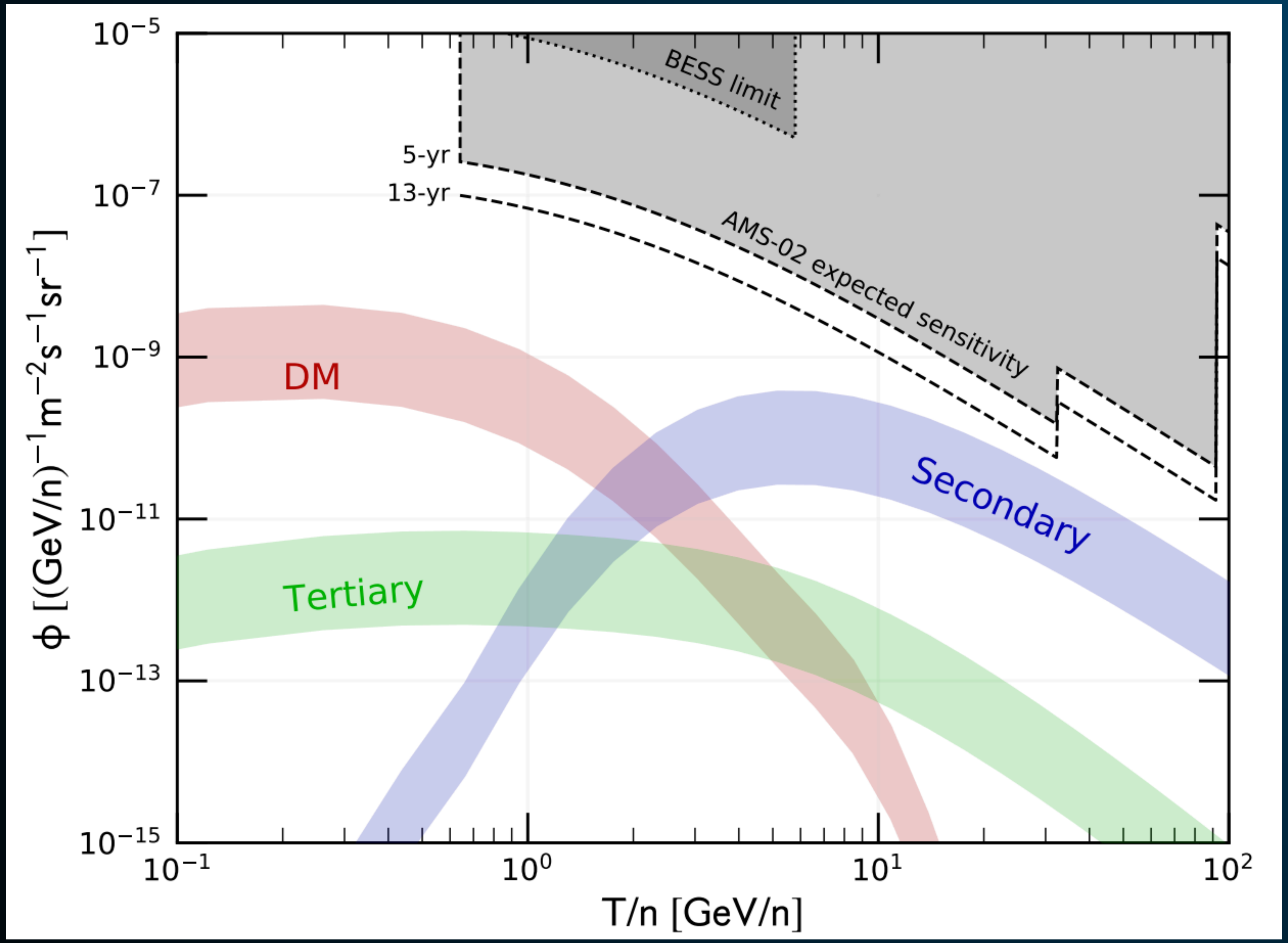
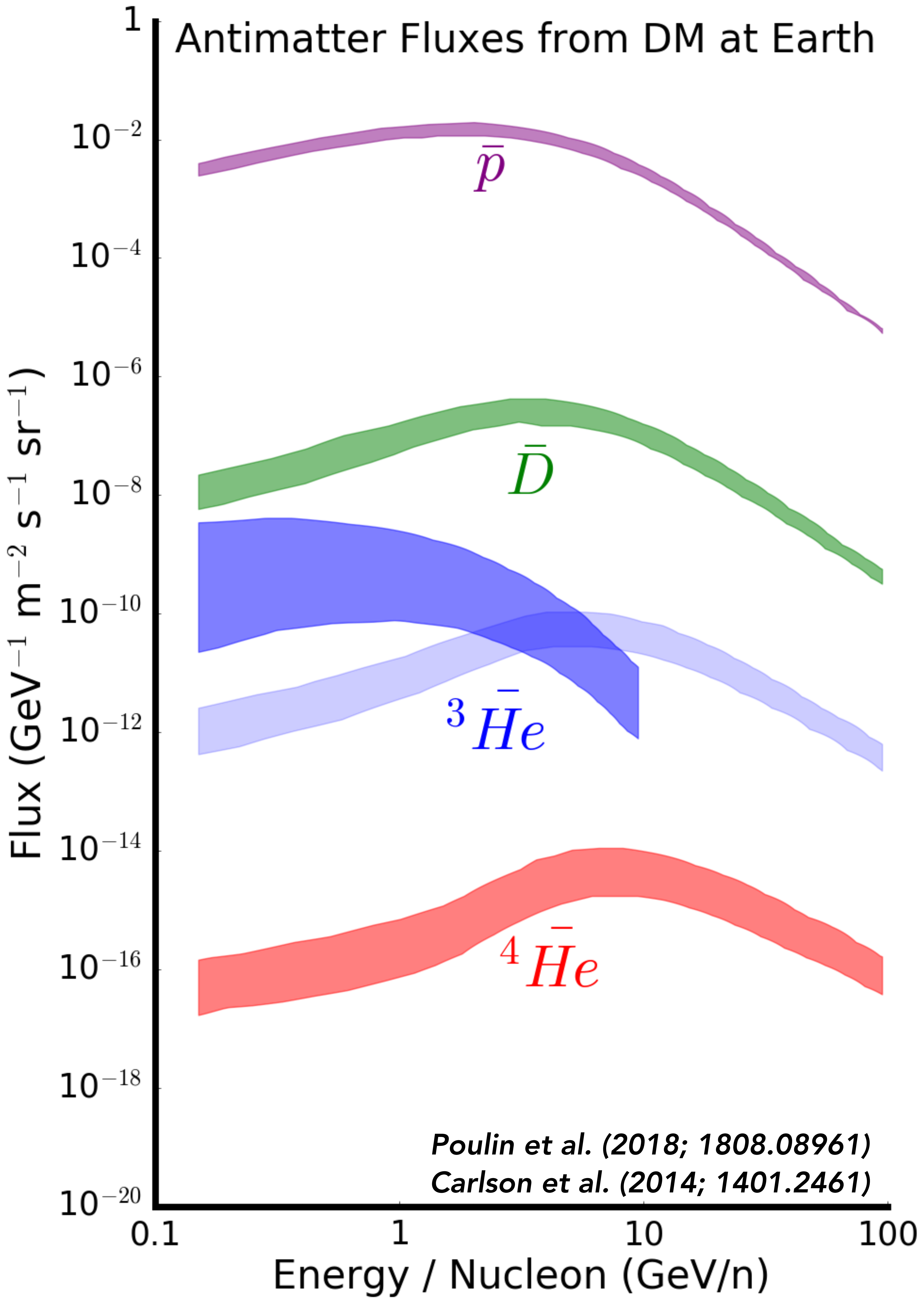
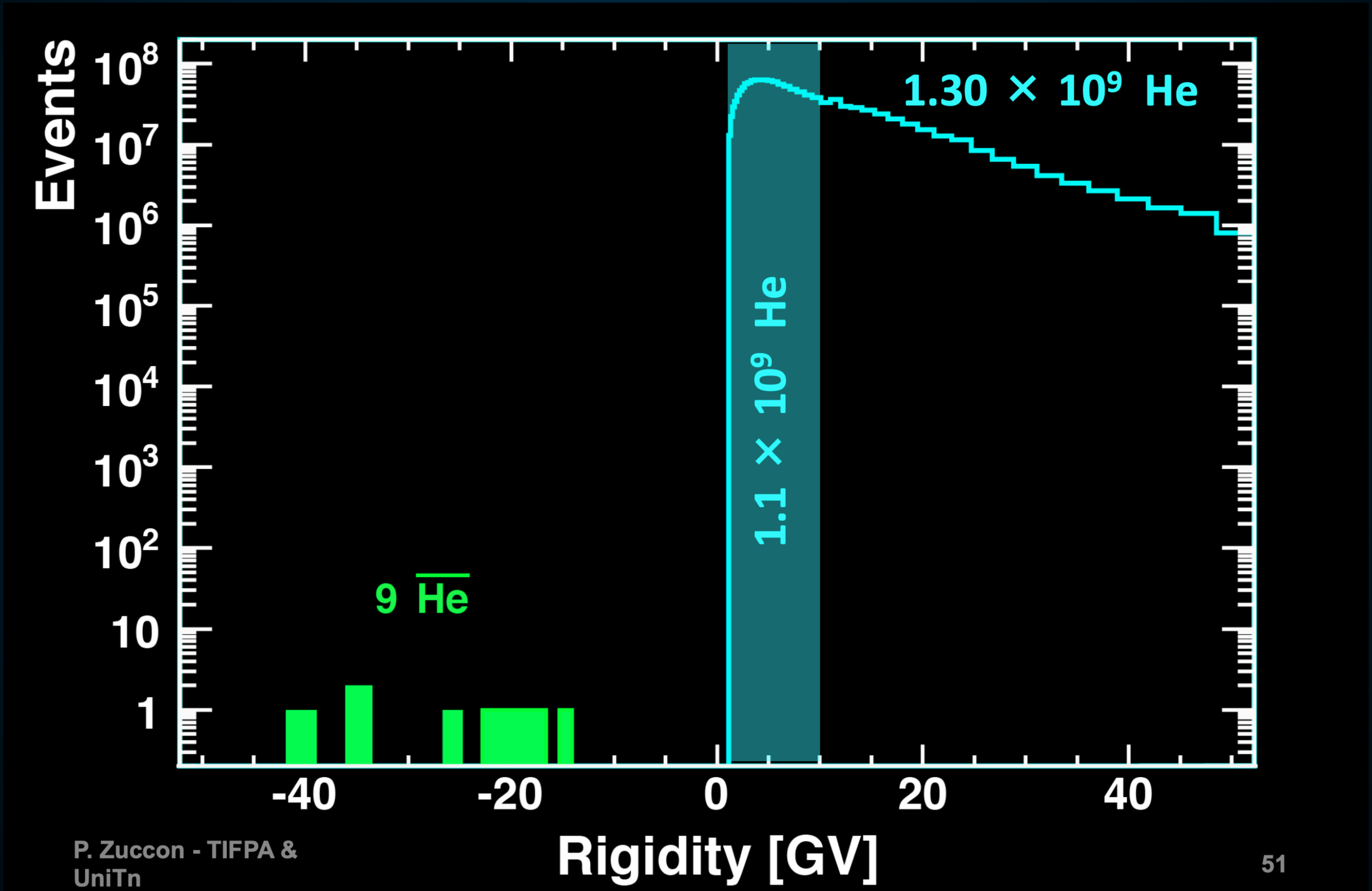


Korsmeier (2017; 1711.08465)



However the Rigidity of these Antihelium Events is High



Idea 2: A New Method for Producing Antihelium

Dark Matter Annihilation Can Produce a Detectable Antihelium Flux through $\bar{\Lambda}_b$ Decays

Martin Wolfgang Winkler[1, *] and Tim Linden[1, †]

[1] *Stockholm University and The Oskar Klein Centre for Cosmoparticle Physics, Alba Nova, 10691 Stockholm, Sweden*

Recent observations by the Alpha Magnetic Spectrometer (AMS-02) have tentatively detected a handful of cosmic-ray antihelium events. Such events have long been considered as smoking-gun evidence for new physics, because astrophysical antihelium production is expected to be negligible. However, the dark-matter-induced antihelium flux is also expected to fall below current sensitivities, particularly in light of existing antiproton constraints. Here, we demonstrate that a previously neglected standard model process — the production of antihelium through the displaced-vertex decay of $\bar{\Lambda}_b$-baryons — can significantly boost the dark matter induced antihelium flux. This process can triple the standard prompt-production of antihelium, and more importantly, entirely dominate the production of the high-energy antihelium nuclei reported by AMS-02.

ep-ph] 29 Jun 2020

I. INTRODUCTION

The detection of massive cosmic-ray antinuclei has long been considered a holy grail in searches for WIMP dark matter [1, 2]. Primary cosmic-rays from astrophysical sources are matter-dominated, accelerated by nearby supernova, pulsars, and other extreme objects. The secondary cosmic-rays produced by the hadronic interactions of primary cosmic-rays can include an antinuclei component, but the flux is highly suppressed by baryon number conservation and kinematic constraints [3, 4]. Dark matter annihilation, on the other hand, occurs within the rest frame of the Milky Way and produces equal baryon and antibaryon fluxes [1, 5–7]

In this *letter*, we challenge the current understanding that standard dark matter annihilation models cannot produce a measurable antihelium flux. Our analysis examines a known, and potentially dominant, antinuclei production mode which has been neglected by previous literature – the production of antihelium through the off-vertex decays of the $\bar{\Lambda}_b$. Such bottom baryons are generically produced in dark matter annihilation channels involving b quarks. Their decays efficiently produce heavy antinuclei due to their antibaryon number and 5.6 GeV rest-mass, which effectively decays to multi-nucleon states with small relative momenta. Intriguingly, because any $^3\overline{\text{He}}$ produced by $\bar{\Lambda}_b$ inherits its boost factor, these nuclei can obtain the large center-of-mass momenta necessary to fit AMS-02 data [13].

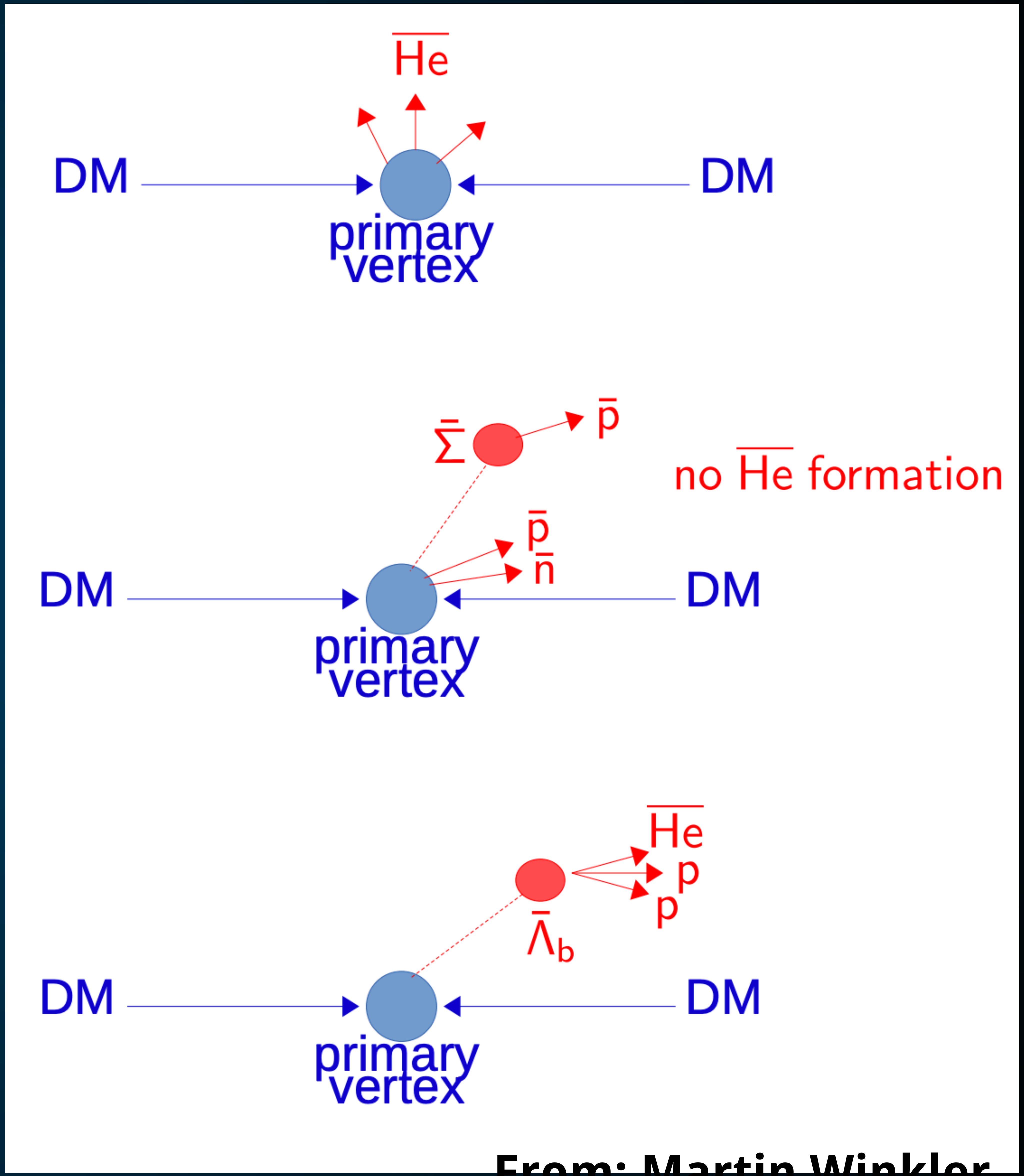
A Standard Model Resonance to Enhance Antihelium

Previous analyses have missed the (potentially) dominant contribution to anti-Helium production.

$\overline{\Lambda_b}$ **has correct parameters to produce** $^3\overline{He}$**:**

- **Antibaryon number of 1**
  - **Mass: 5.6 GeV** $(\bar{p}, \bar{p}, \bar{n}, p, p)$
  - **Or:** $\bar{p}, \bar{n}, \bar{n}, p, p$ **because** $^3H \rightarrow {}^3He$

$$R \propto p_0^{3(A-1)} \qquad R \propto \exp[-(\mathrm{p_i} - \mathrm{p_f})]$$

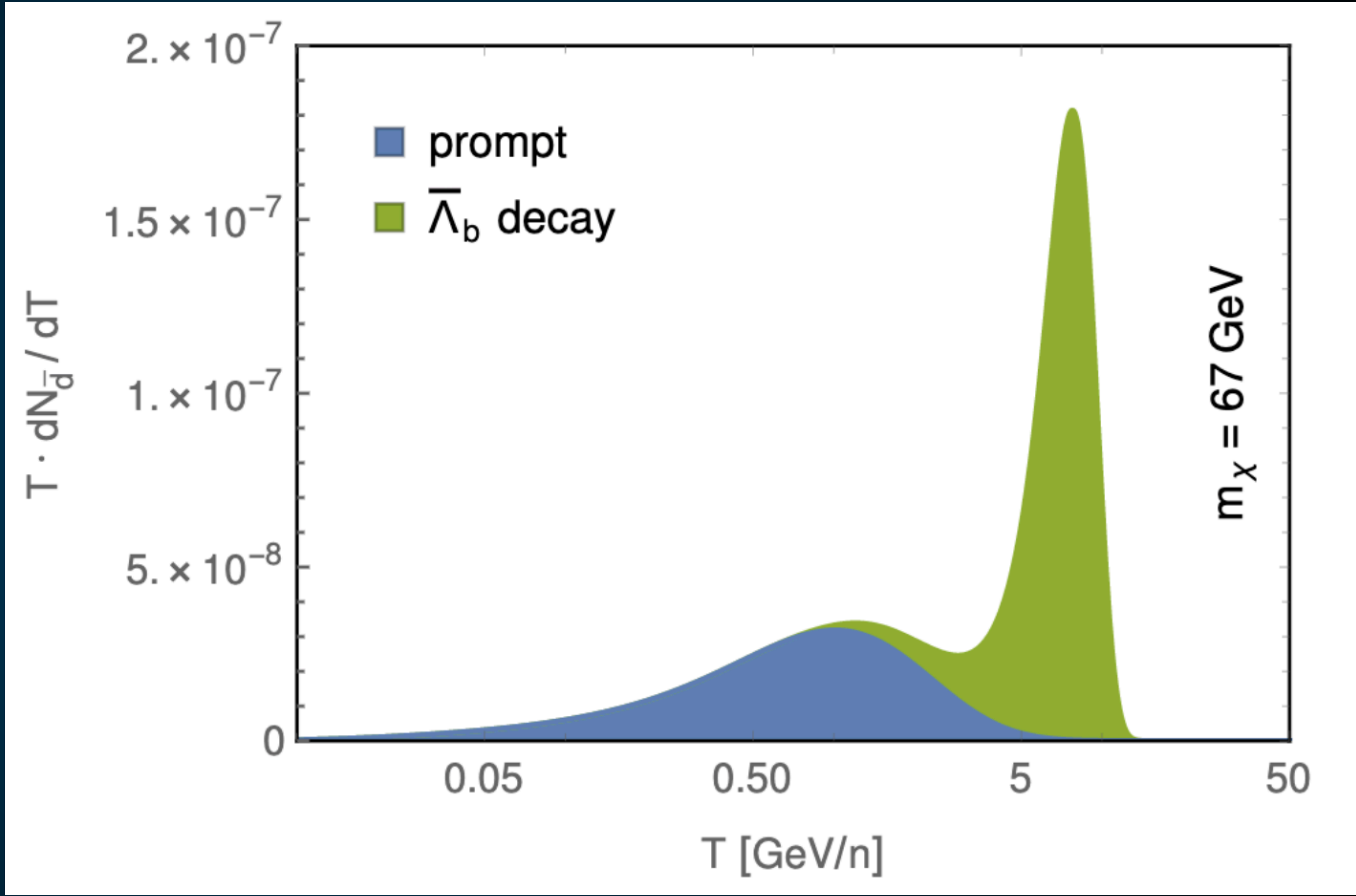


From: Martin Winkler

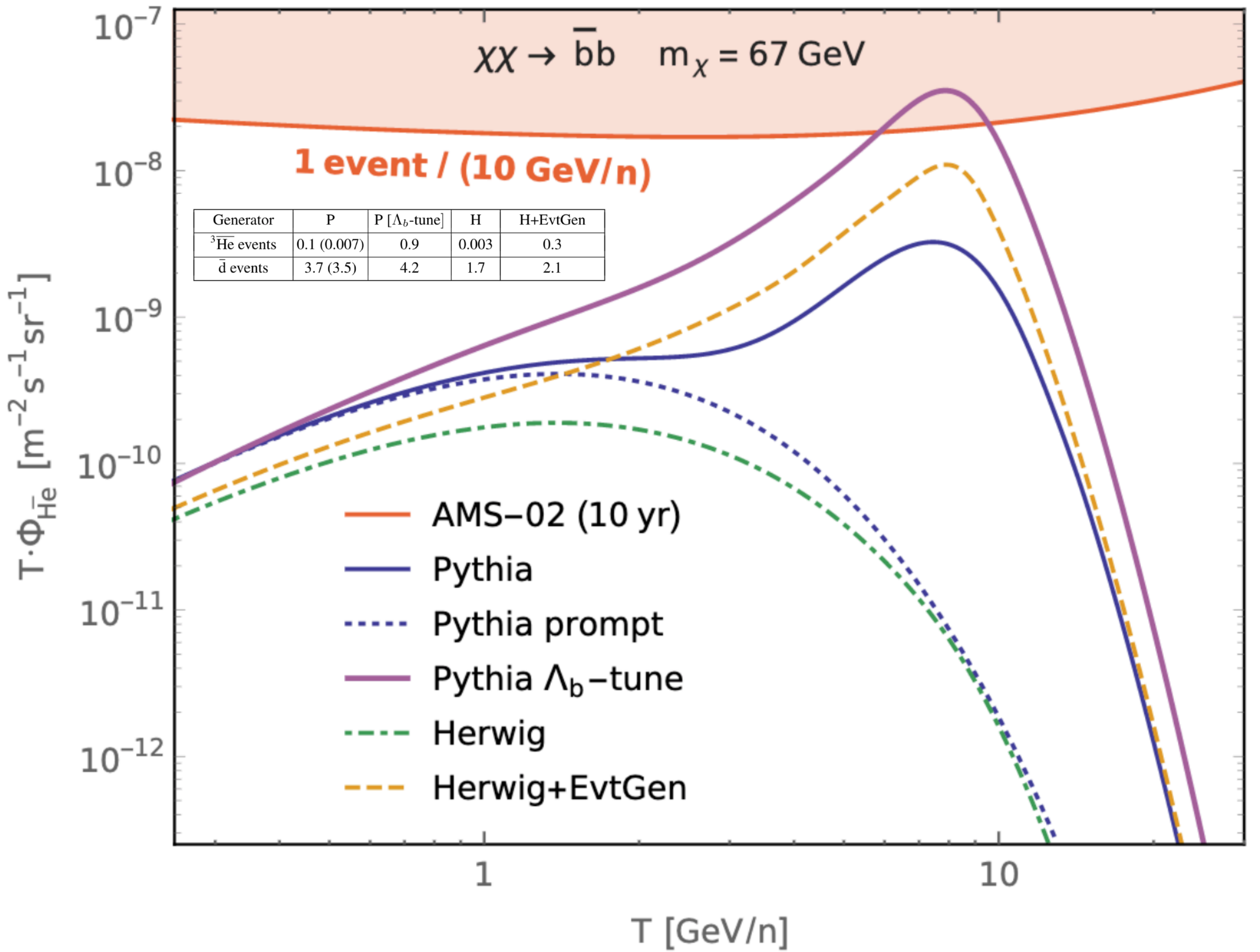
A High-Momentum Bump!

Can produce a significant enhancement of the total anti helium flux.

Moreover, the enhancement is at high-energies - matching the data.



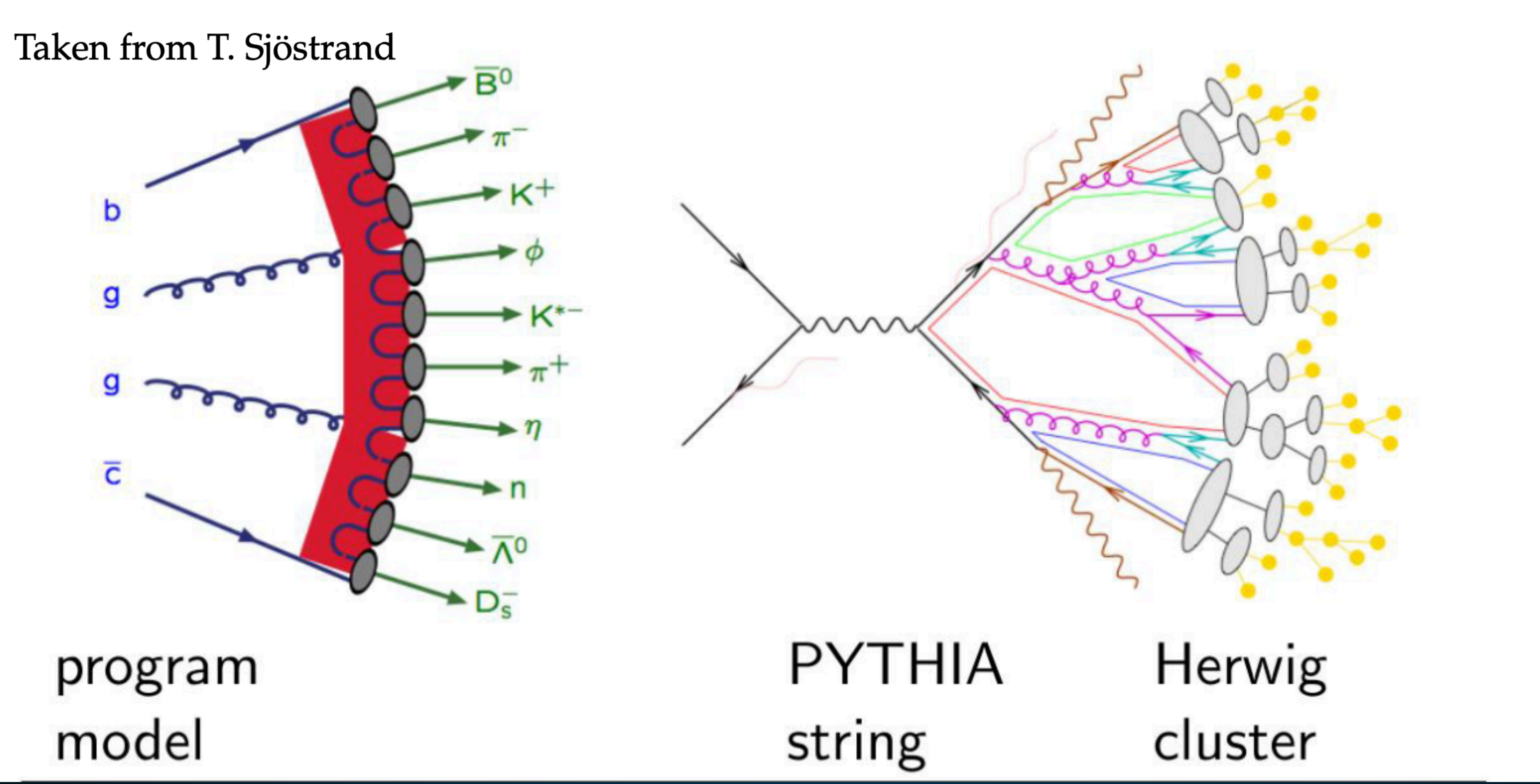
Winkler & Linden (2020; 2020.16251)



Generator	P	P [Λ_b-tune]	H	H+EvtGen
$^3\overline{He}$ events	0.1 (0.007)	0.9	0.003	0.3
$\bar{d}$ events	3.7 (3.5)	4.2	1.7	2.1

Uncertainties in the Rate

$\overline{\Lambda_b} \to {}^3\text{He}$ rate



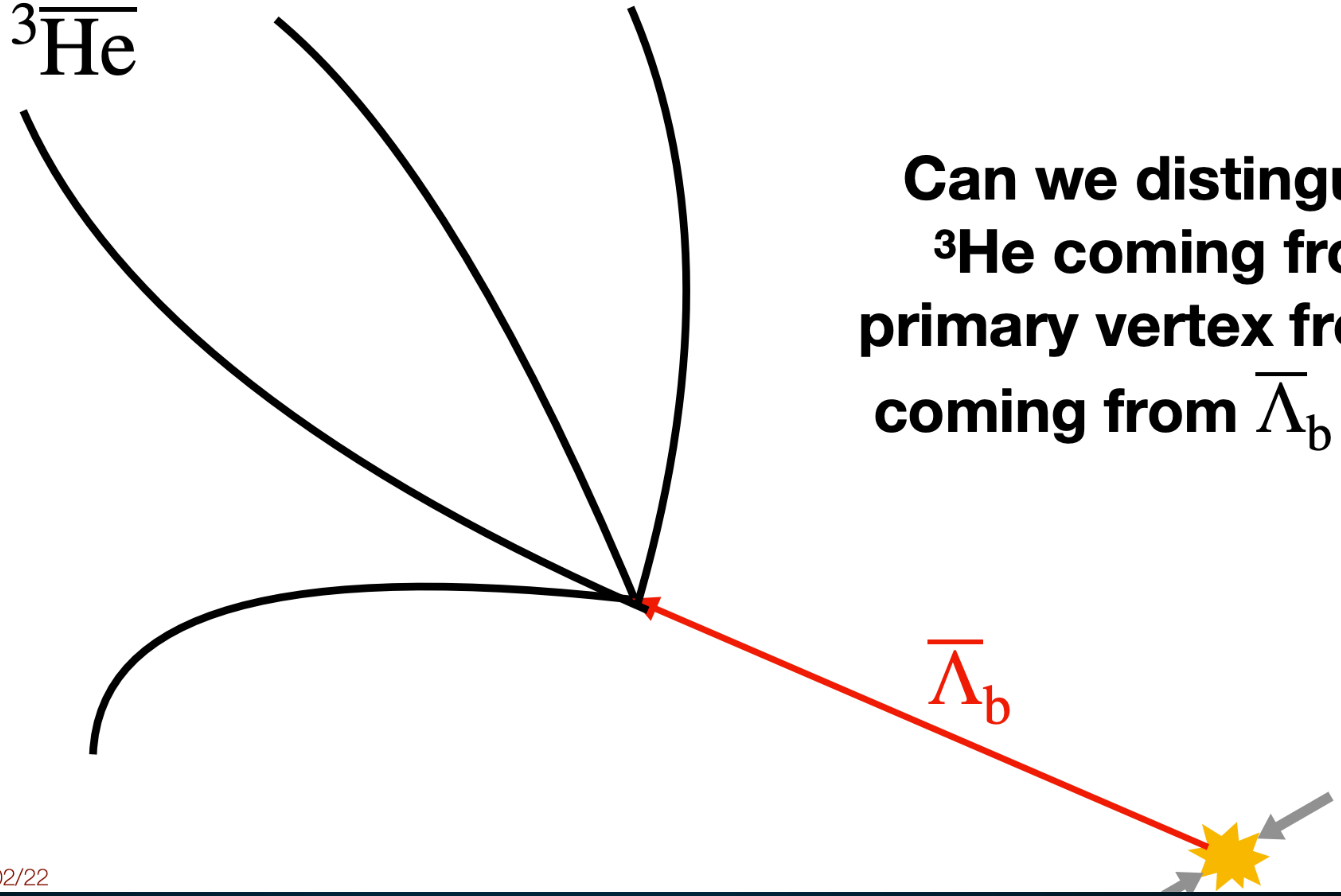
Pythia:

$$P(\overline{\Lambda_b} \to {}^3\overline{\text{He}} + X) \sim 10^{-6}$$

Herwig:

$$P(\overline{\Lambda_b} \to {}^3\overline{\text{He}} + X) \sim 10^{-9}$$

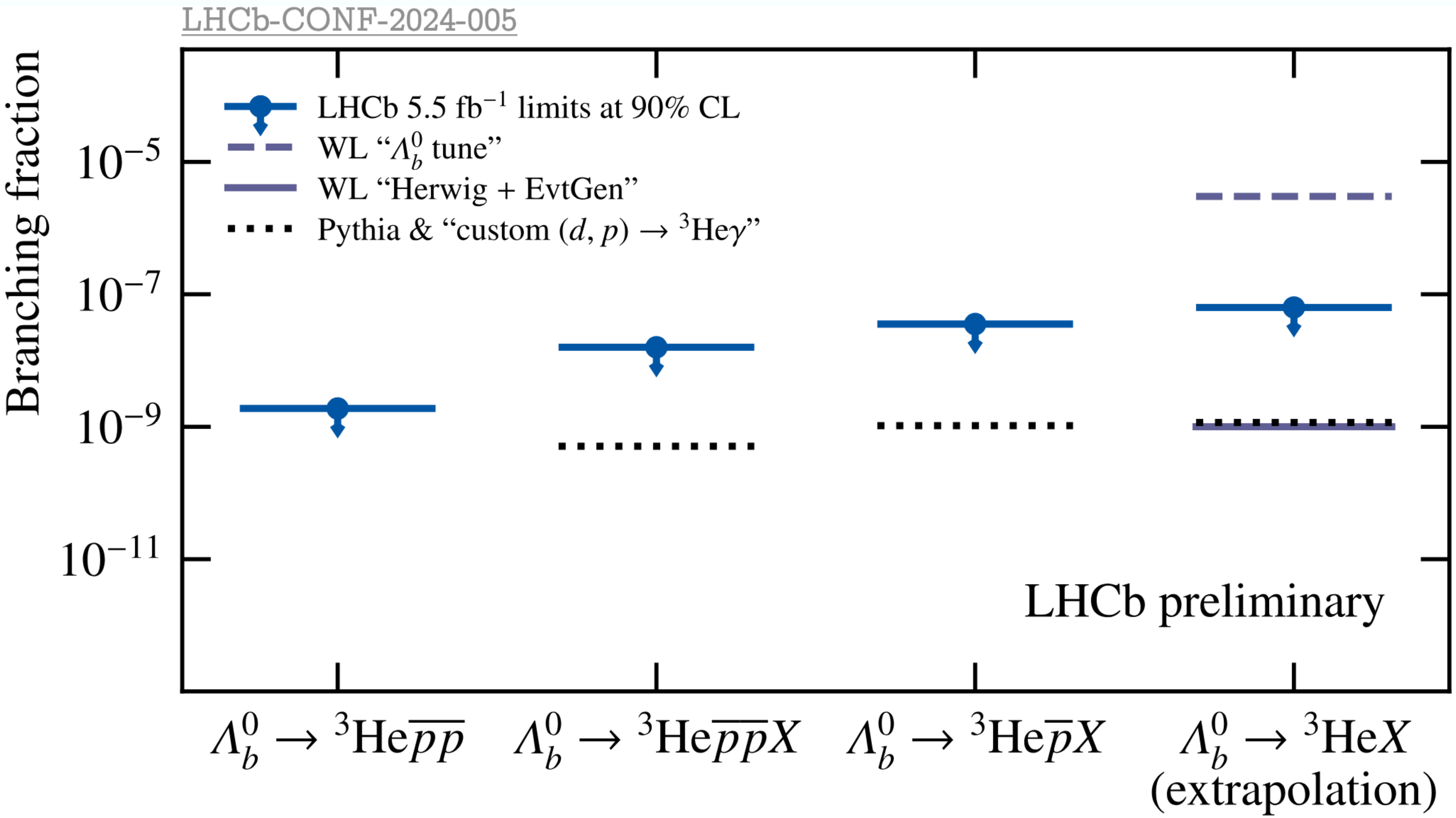
Can We Find this At Particle Accelerators?



Can we distinguish the ^{3}He coming from the primary vertex from those coming from $\overline{\Lambda}_b$ decays?

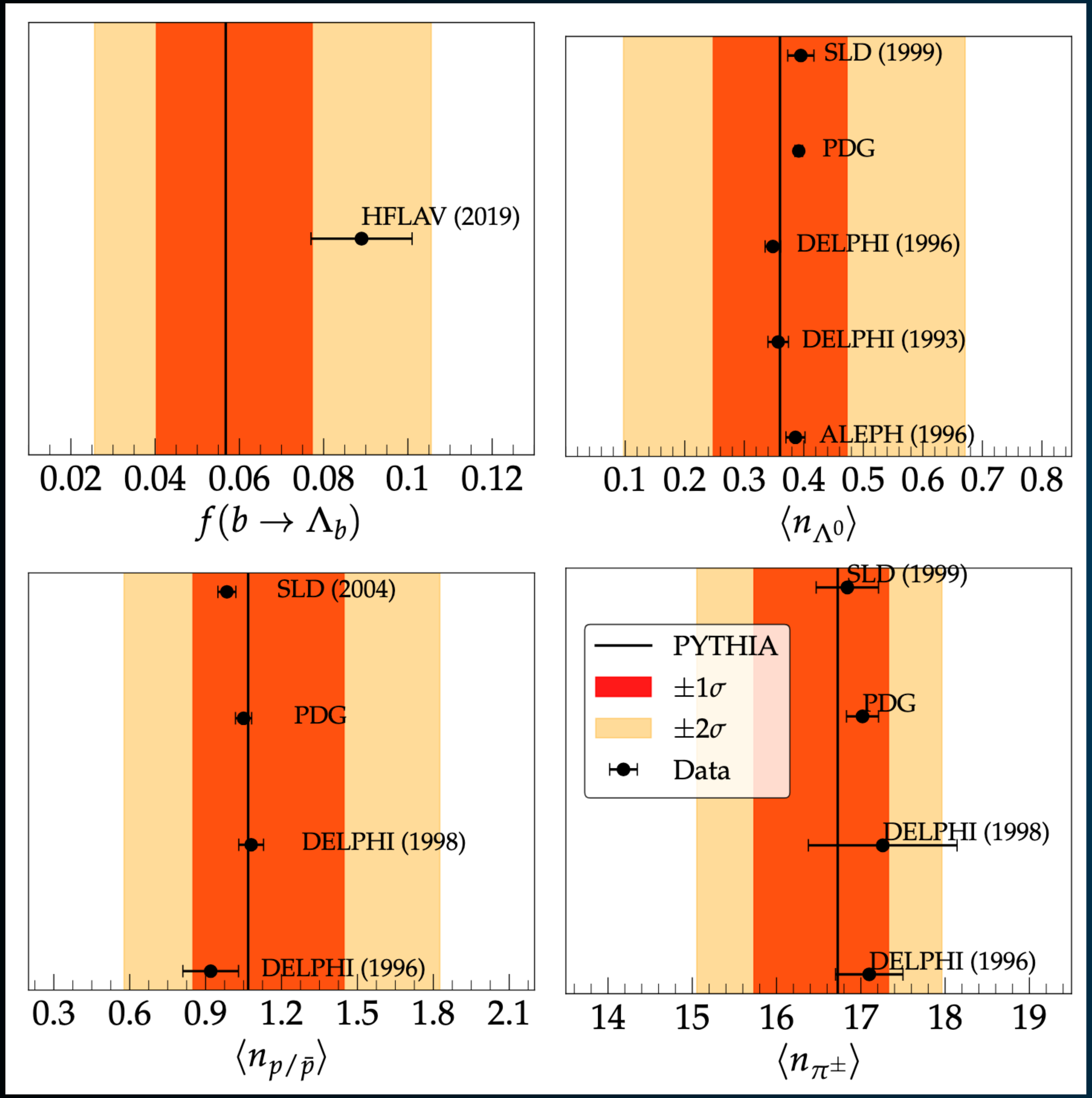
$c\tau$ ~ 450 µm

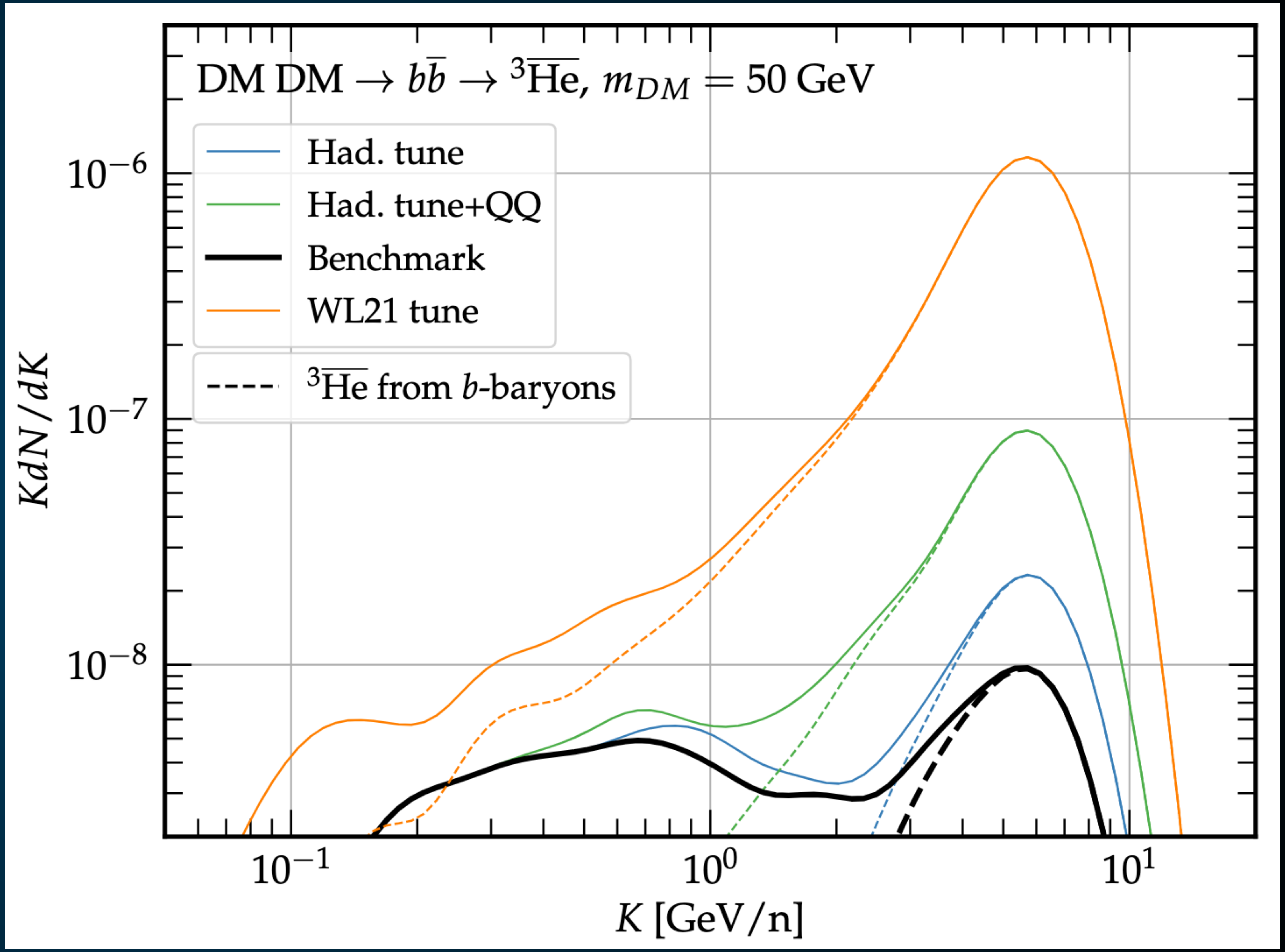
Current status



These are the first results on (anti)helium production in (anti)Λ_b decays.
They significantly restrict abundance of ${}^3\overline{\mathrm{He}}$ in cosmic rays.

Uncertainties in the Rate





There are tunes of Pythia that decrease the expected anti-helium flux.

Di Mauro et al. (2504.07172)

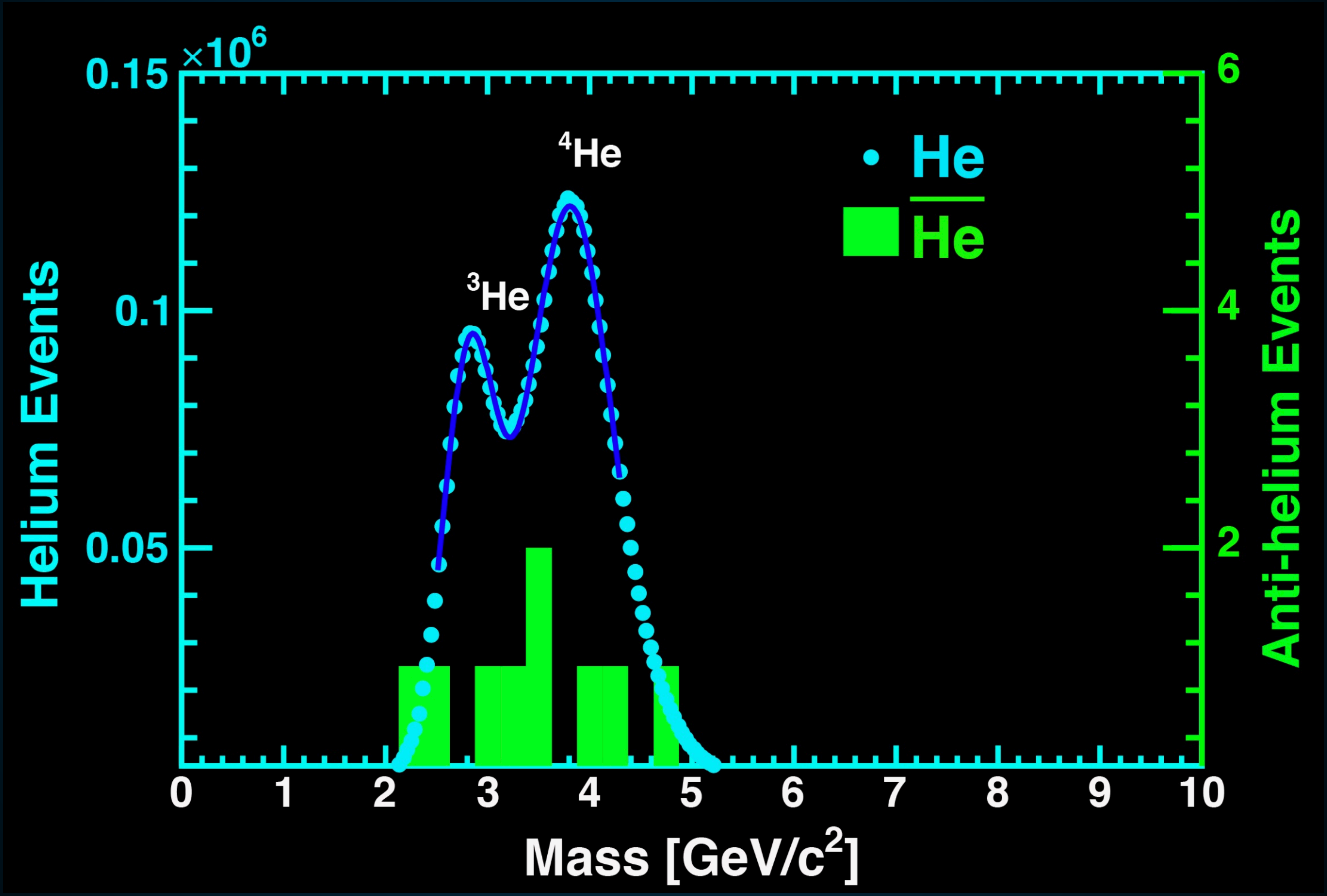
Some Caveats

1.) There is a factor of two offset, because tritium decays to ${}^3\mathrm{He}$ in space.

- This can potentially be larger, because $\overline{p} + \overline{n} + \overline{n} + p + n$ has smaller kinetic energy.
- Also: See Isospin Asymmetry Slides from Thomas Pöschl

2.) No searches for $\overline{{}^3\mathrm{H}} + n + n + \pi^+$. This could dominate, for example, if the proton and $\overline{{}^3\mathrm{He}}$ quickly re-annihilate due to Coulomb attraction.

Problem: Are We Actually Observing Antihelium 4?

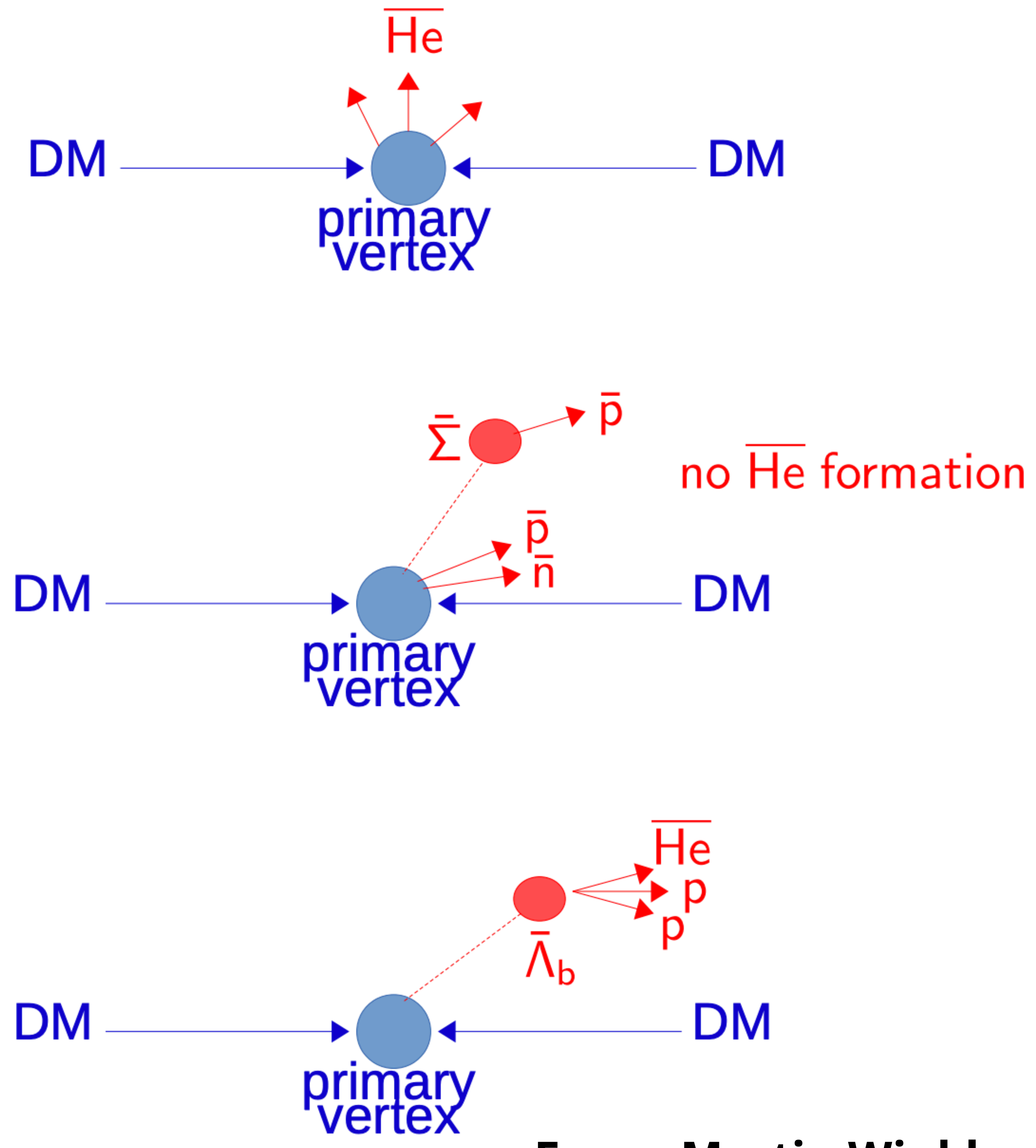


Cannot Enhance Antihelium-4 with Λ_b

$\overline{\Lambda}_b$ has correct parameters to produce $^3\overline{\mathrm{He}}$:

- Antibaryon number of 1
- Mass: 5.6 GeV

Too light to produce $^4\overline{\mathrm{He}}$!



From: Martin Winkler

Strongly Coupled Dark Sectors

Cosmic Ray Antihelium from a Strongly Coupled Dark Sector

Martin Wolfgang Winkler,[1,2,*] Pedro De La Torre Luque,[2,†] and Tim Linden[2,‡]

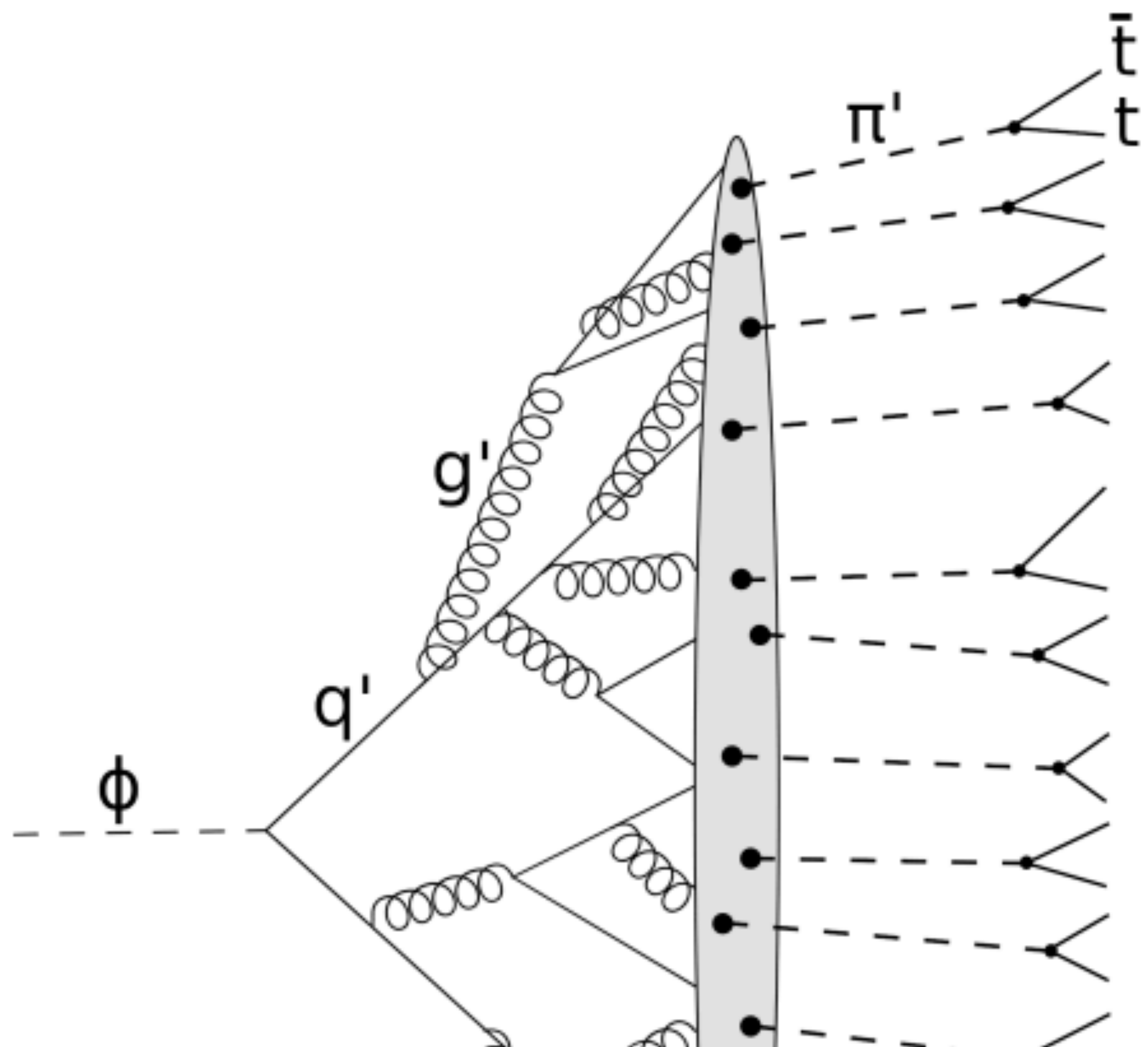
[1] *Department of Physics, The University of Texas at Austin, Austin, 78712 TX, USA*

[2] *The Oskar Klein Centre, Department of Physics, Stockholm University, AlbaNova, SE-10691 Stockholm, Sweden*

Standard Model extensions with a strongly coupled dark sector can induce high-multiplicity states of soft quarks. Such final states trigger extremely efficient antinucleus formation. We show that dark matter annihilation or decay into a strongly coupled sector can dramatically enhance the cosmic-ray antinuclei flux – by six orders of magnitude in the case of $^4\overline{\mathrm{He}}$. In this work, we argue that the tentative $^3\overline{\mathrm{He}}$ and $^4\overline{\mathrm{He}}$ events reported by the AMS-02 collaboration could be the first sign of a strongly coupled dark sector observed in nature.

I. INTRODUCTION

Cosmic-ray (CR) antinuclei are among the most promising targets in the indirect search for particle dark matter (DM). While the formation of antinuclei by DM annihilation or decay is strongly suppressed compared to *e.g.* gamma rays, the astrophysical antinuclei backgrounds – which arise from interactions of cosmic ray protons and helium with the interstellar gas – are extremely low. Therefore, the unambiguous discovery of even a single cosmic-ray antinucleus could provide smoking-gun evidence for particle DM [1, 2].



Strongly Coupled Dark Sectors

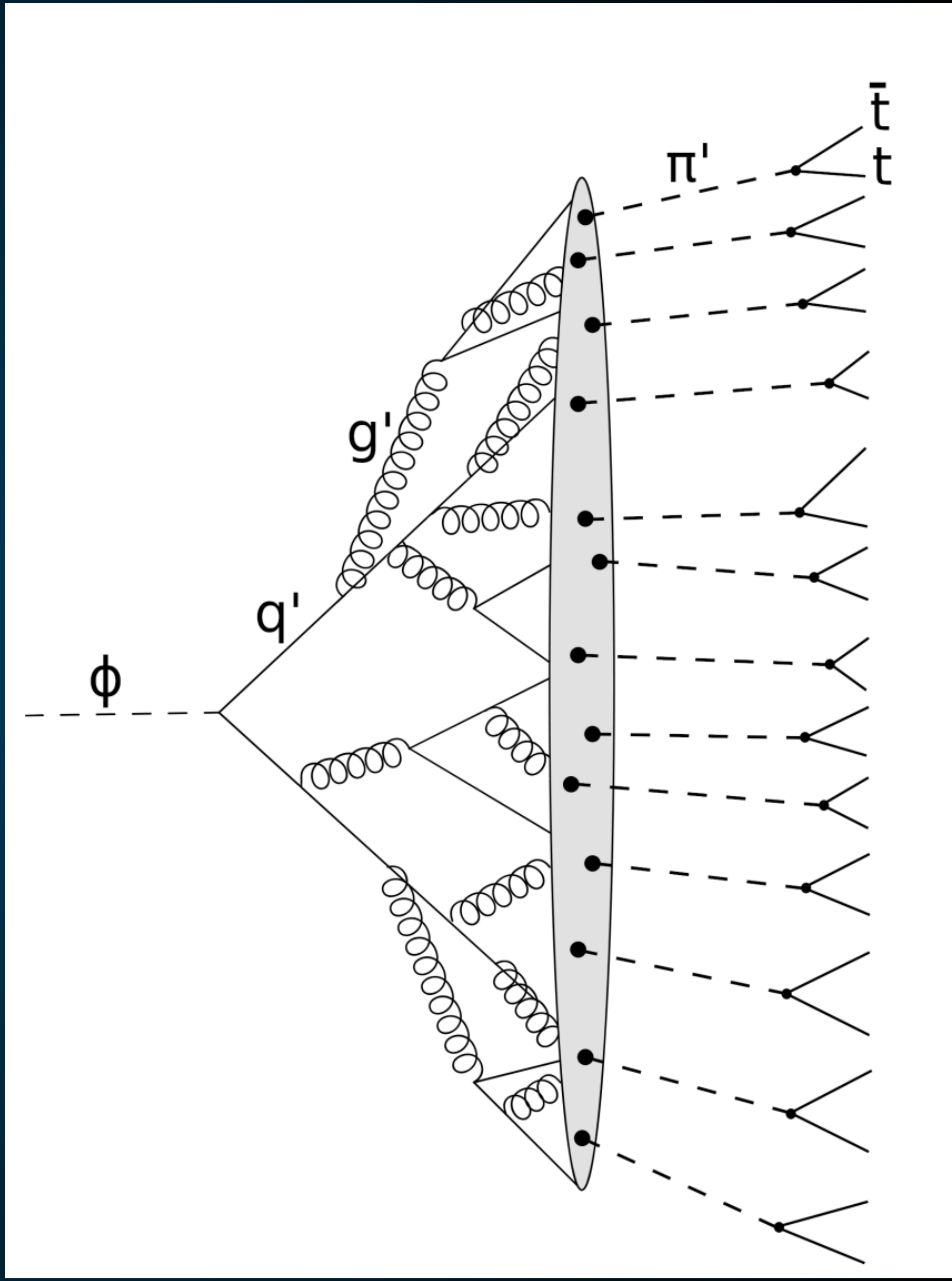
Just make a ton of quarks.

The production of heavy nuclei scales strongly with the number of quarks in the final state.

In QCD, a single 100 GeV annihilation produces O(100) pions

The dark matter model looks like a dark version of QCD.

$$\mathcal{L} \supset -\frac{1}{2}\,\mathrm{Tr}\,G'_{\mu\nu}G'^{\mu\nu} - \bar{q}'(i\not{D} - m_{q'})q'$$



Strongly Coupled Dark Sectors

The dark pions need to be very heavy — so the dark matter also has to be very heavy.

For annihilating dark matter — we are limited by unitarity.

For decaying dark matter, we are not.

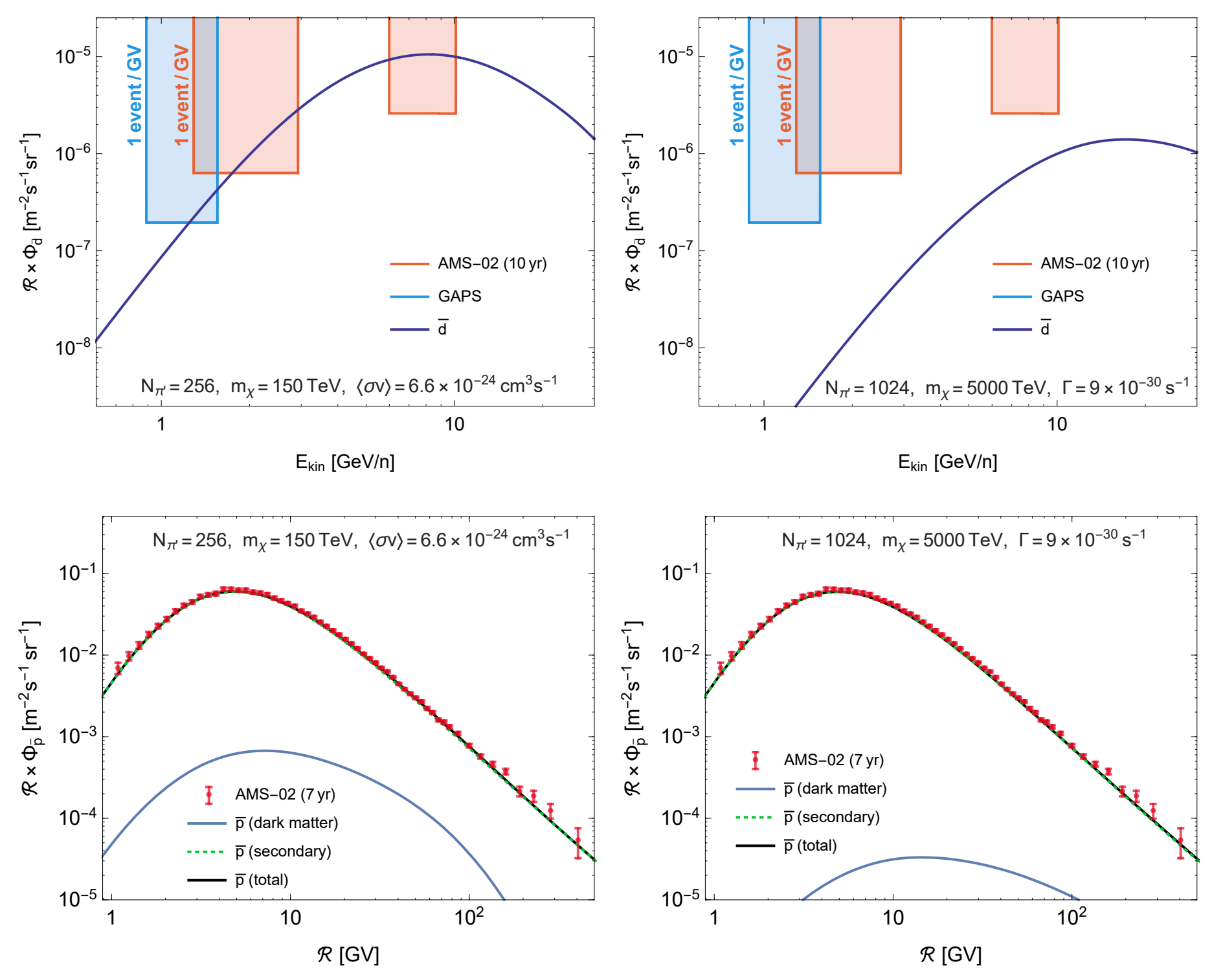
DM type	Annihilating	Decaying
Input Parameters		
m_χ [TeV]	150	5000
m_ϕ [TeV]	50.4	375
$m_{\pi'}$ [GeV]	380	700
$N_{\pi'}$	256	1024
$\langle\sigma v\rangle$ [$\mathrm{cm^3 s^{-1}}$]	6.6×10^{-24}	–
Γ [$\mathrm{s^{-1}}$]	–	9×10^{-30}
Antinuclei Events at AMS-02		
$^3\overline{\mathrm{He}}$	15.6	20.3
$^4\overline{\mathrm{He}}$	1.0	3.1
$\bar{\mathrm{d}}$	19.3	1.2
Antinuclei Events at GAPS		
$\bar{\mathrm{d}}$	0.7	0

Strongly Coupled Dark Sectors

This significantly boosts the anti helium production rate:

By a factor of n^9 for $^3\overline{He}$ and n^{12} for $^4\overline{He}$

Strongly Coupled Dark Sectors



Increasingly Unlikely Methods for Producing Antihelium

Coalescence - Essentially predicts the number of quarks that will fuse into anti helium

Reacceleration - Important when the particle is very low energy (from coalescence) and also has charge +2 (antihelium specific)

Λ_b - Antiproton decays were already accounted for by standard models, ability to produce antihelium was not known.

Dark Sectors - Antiproton production is small due to heavy DM mass. Antihelium is enhanced (compared to typical rates for heavy dark matter, ~O(10^6).

Where do the *AMS-02* anti-helium events come from?

Vivian Poulin,[1] Pierre Salati,[2] Ilias Cholis,[3,1] Marc Kamionkowski,[1] and Joseph Silk[1,4,5]

[1]*Department of Physics and Astronomy, Johns Hopkins University, Baltimore, MD 21218, USA*
[2]*LAPTh, Université Savoie Mont Blanc & CNRS, 74941 Annecy Cedex, France*
[3]*Department of Physics, Oakland University, Rochester, MI 48309, USA*
[4]*Sorbonne Universités, UPMC Univ. Paris 6 et CNRS, UMR 7095, Institut dAstrophysique de Paris, 98 bis bd Arago, 75014 Paris, France*
[5]*Beecroft Institute of Particle Astrophysics and Cosmology, Department of Physics, University of Oxford, Denys Wilkinson Building, 1 Keble Road, Oxford OX1 3RH, UK*

(Dated: March 26, 2019)

We discuss the origin of the anti-helium-3 and -4 events possibly detected by *AMS-02*. Using up-to-date semi-analytical tools, we show that spallation from primary hydrogen and helium nuclei onto the ISM predicts a $\overline{^{3}\mathrm{He}}$ flux typically one to two orders of magnitude below the sensitivity of *AMS-02* after 5 years, and a $\overline{^{4}\mathrm{He}}$ flux roughly 5 orders of magnitude below the *AMS-02* sensitivity. We argue that dark matter annihilations face similar difficulties in explaining this event. We then entertain the possibility that these events originate from anti-matter-dominated regions in the form of anti-clouds or anti-stars. In the case of anti-clouds, we show how the isotopic ratio of anti-helium nuclei might suggest that BBN has happened in an inhomogeneous manner, resulting in anti-regions with a anti-baryon-to-photon ratio $\bar{\eta} \simeq 10^{-3}\eta$. We discuss properties of these regions, as well as relevant constraints on the presence of anti-clouds in our Galaxy. We present constraints from the survival of anti-clouds in the Milky-Way and in the early Universe, as well as from CMB, gamma-ray and cosmic-ray observations. In particular, these require the anti-clouds to be almost free of normal matter. We also discuss an alternative where anti-domains are dominated by surviving anti-stars. We suggest that part of the unindentified sources in the 3FGL catalog can originate from anti-clouds or anti-stars. *AMS-02* and *GAPS* data could further probe this scenario.

Conclusions

These are non-standard approaches. Even if dark matter is a WIMP, it may not produce antihelium.

However, if antihelium is detected, these are among the most reasonable methods for producing such an exotic particle.

All of these avenues are experimentally testable with upcoming colliders.

Uncertainties in the Rate

1.) $\overline{\Lambda}_b$ production rate

Comment on "Dark Matter Annihilation Can Produce a Detectable Antihelium Flux through $\bar{\Lambda}_b$ Decays"

M. Kachelrieß[1], S. Ostapchenko[2], and J. Tjemsland[1]
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In a recent Letter, Winkler and Linden [1] (hereafter WL21) suggested that a previously neglected standard model process, namely the production of antihelium-3 nuclei through decays of $\bar{\Lambda}_b$ baryons, can significantly boost the flux of antihelium-3, induced by annihilations or decays of dark matter. This suggestion uses the fact that dark matter particles will annihilate typically into the heaviest quark–anti-quark pair, i.e. $\bar{b}b$ pairs, if the particle is a Majorana fermion and its mass is below the mass of the standard model gauge bosons [2]. These (anti-) b quarks will in turn hadronise and form (anti-) b-mesons and (anti-) b-baryons which then decay weakly. As pointed out by WL21, the $\bar{\Lambda}_b$ baryon is especially suited for the production of antihelium-3 through a coalescence process, because its rest mass of 5.6 GeV is not much above the rest mass of 5 (anti)-nucleons. As a result of the small relative momenta of these nucleons, the production of antihelium-3 via coalescence is enhanced in $\bar{\Lambda}_b$ decays.

of the coalescence momentum, since this change affects all types of processes involving baryon and meson production. As an example, one can consider (anti-) proton production in electron-positron annihilations, $e^+e^- \to \bar{p}pX$. For a change of `probQQtoQ` from the default value 0.09 to 0.24—which is the value reproducing the value of the branching ratio $b \to \Lambda_b = 0.1$ chosen in WL21—the resulting proton multiplicity is compared in Table 1 to measurements. For instance at $\sqrt{s} = 91\,\text{GeV}$, the predicted proton multiplicity in the "$\bar{\Lambda}_b$ tune" is 33σ away from the one measured [3]. For comparison, the standard settings in Pythia predict a Λ_b multiplicity in electron-positron annihilations at the Z-resonance of 0.016, which is less than $1\,\sigma$ away from the value 0.031 ± 0.016 given in Ref. [3]. As an example for the effects of a changed diquark formation parameter on pp collisions, we show in Table 2 the integrated yield at mid-rapidity, $\mathrm{d}N/\mathrm{d}y\,|_{|y|<0.5}$, of protons, kaons and pions measured by ALICE at LHC at $\sqrt{s} = 7\,\text{TeV}$ [4]. Note also that the

Claims:

1.) Changing probQQtoQ is not needed because previous data was within uncertainties.

2.) Changing probQQtoQ breaks models of light baryon formation (at 33σ in certain LEP cross-sections)

Uncertainties in the Rate

1.) $\overline{\Lambda}_b$ production rate

Response to Comment on "Dark Matter Annihilation Can Produce a Detectable Antihelium Flux through $\bar{\Lambda}_b$ Decays"

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In a recent paper we showed that the decay of intermediate $\bar{\Lambda}_b$ baryons can dramatically enhance the antihelium flux from dark matter annihilation. Our antihelium predictions were derived using several implementations of the Pythia and Herwig event generators which were calibrated to existing data on antideuteron and antihelium formation. Kachelrieß et al. have argued for a smaller antihelium flux compared to our most optimistic Monte Carlo model. However, we show that the arguments by Kachelrieß et al. are either incorrect or irrelevant for antihelium formation. Thus, the results of our original paper remain unchanged.

hep-ph] 31 May 2021

Summary of Original Paper – The AMS-02 cosmic-ray experiment has tentatively detected a handful of cosmic-ray antihelium events [1]. This observation is puzzling, since neither astrophysical processes nor dark matter annihilation were thought to produce a detectable antihelium flux. In recent work [2], we demonstrated that a previously neglected standard-model process – the production of antihelium nuclei through the decay of intermediate $\bar{\Lambda}_b$ baryons – could dramatically enhance the antihelium production rate from dark matter annihilation. The key insight is that $\bar{\Lambda}_b$, due to its antibaryon number and 5.6 GeV rest-mass, efficiently decays to multi-antibaryons states with small relative momentum which coalesce into antihelium.

The antihelium flux from $\bar{\Lambda}_b$ decay depends on (1) the rate of $\bar{\Lambda}_b$ baryon production in dark matter annihilations to

As we will show below, these statements are either incorrect or inapplicable to our study. Moreover, we stress that even if the analysis of KOT21 were entirely correct, these criticisms amount to only a factor of ∼3 adjustment in a novel factor of ∼100 effect first pointed out in our original paper.

First, statement (1) is incorrect. The fraction $f(b \to \Lambda_b)$ has been determined by a variety of LEP [4–6], Tevatron [7] and LHC [8] measurements. Several of the most relevant results are listed in Table I, together with the predicted rate from both default Pythia models and our Λ_b-tune model. Standard Pythia implementations predict a Λ_b-production rate that falls below results from several independent experiments and channels at a combined significance exceeding 10σ. Our Λ_b-tune model, on the other hand, provides a good fit to LEP data. It still underpredicts the data from hadron colliders, signal-

Claims:

1.) Untrue: Errors in Pythia Models were above 10σ.

2.) True - but irrelevant - if you want to get b-quark physics right, it is reasonable to tune the model to b-quark physics.

An Astrophysical Solution: Move the Excess to High Energies

1.) Changing the coalescence model primarily affects the Helium yield when the total center of mass energy is small.

2.) Very good for predicted rates with GAPS, or low-energy AMS-02 observations.

3.) But AMS-02 antihelium are (generally reported) at energies of ~10 GeV/n.

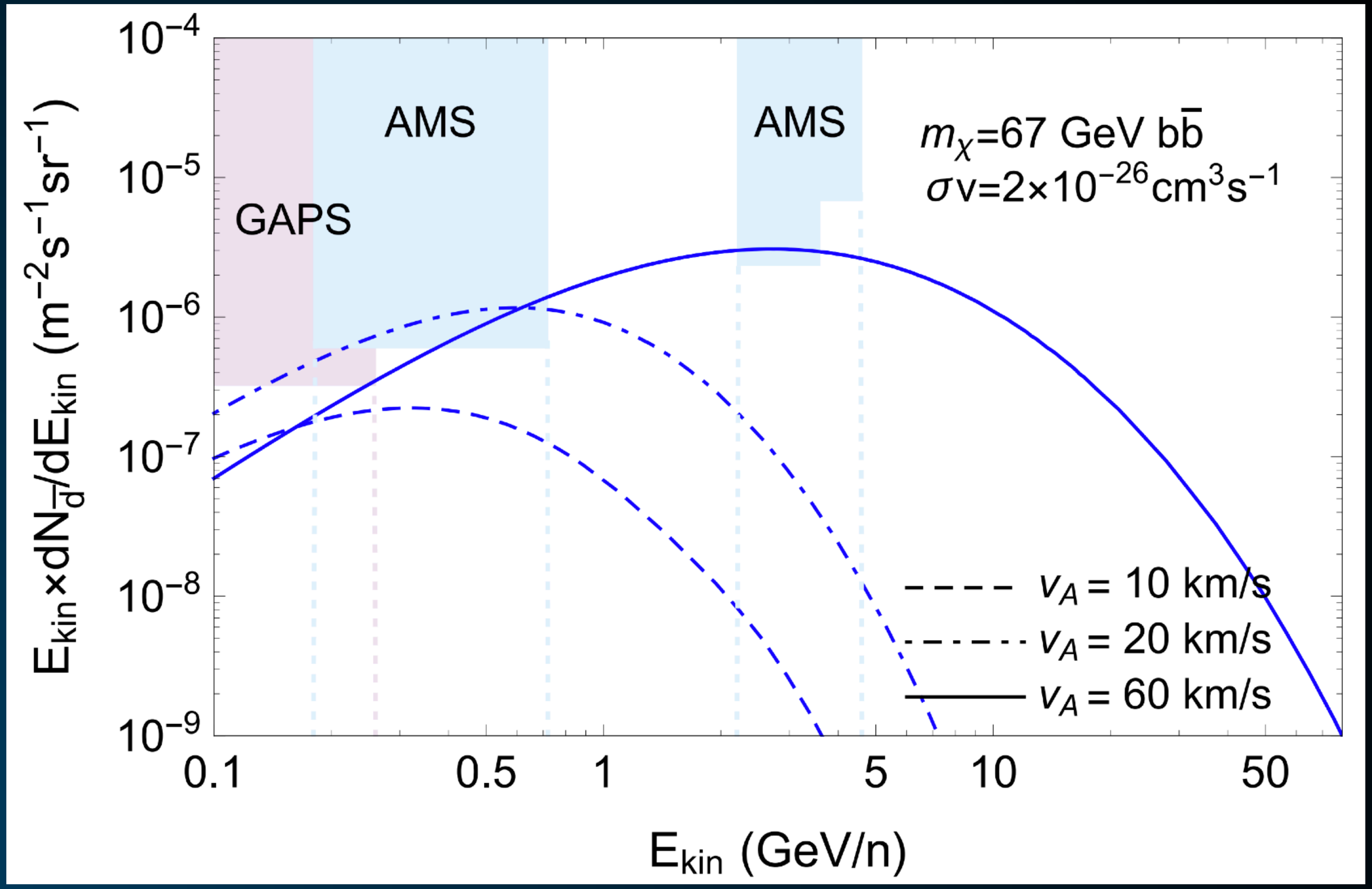
Astrophysical Enhancements!

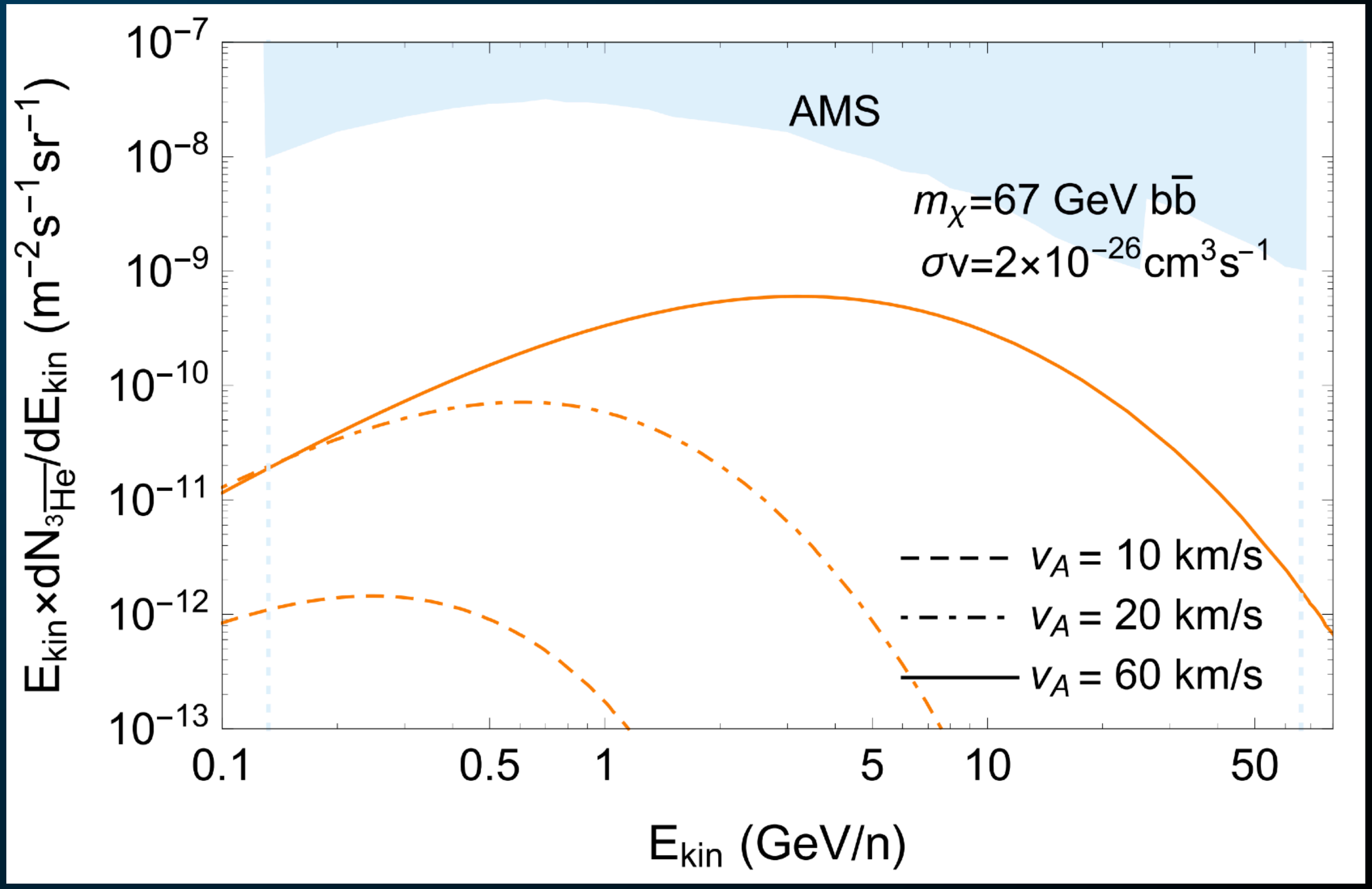
The current event rates depend on the detector sensitivity to anti-Helium.

We lose many events because most anti-He are produced at energies that are too small to be detected.

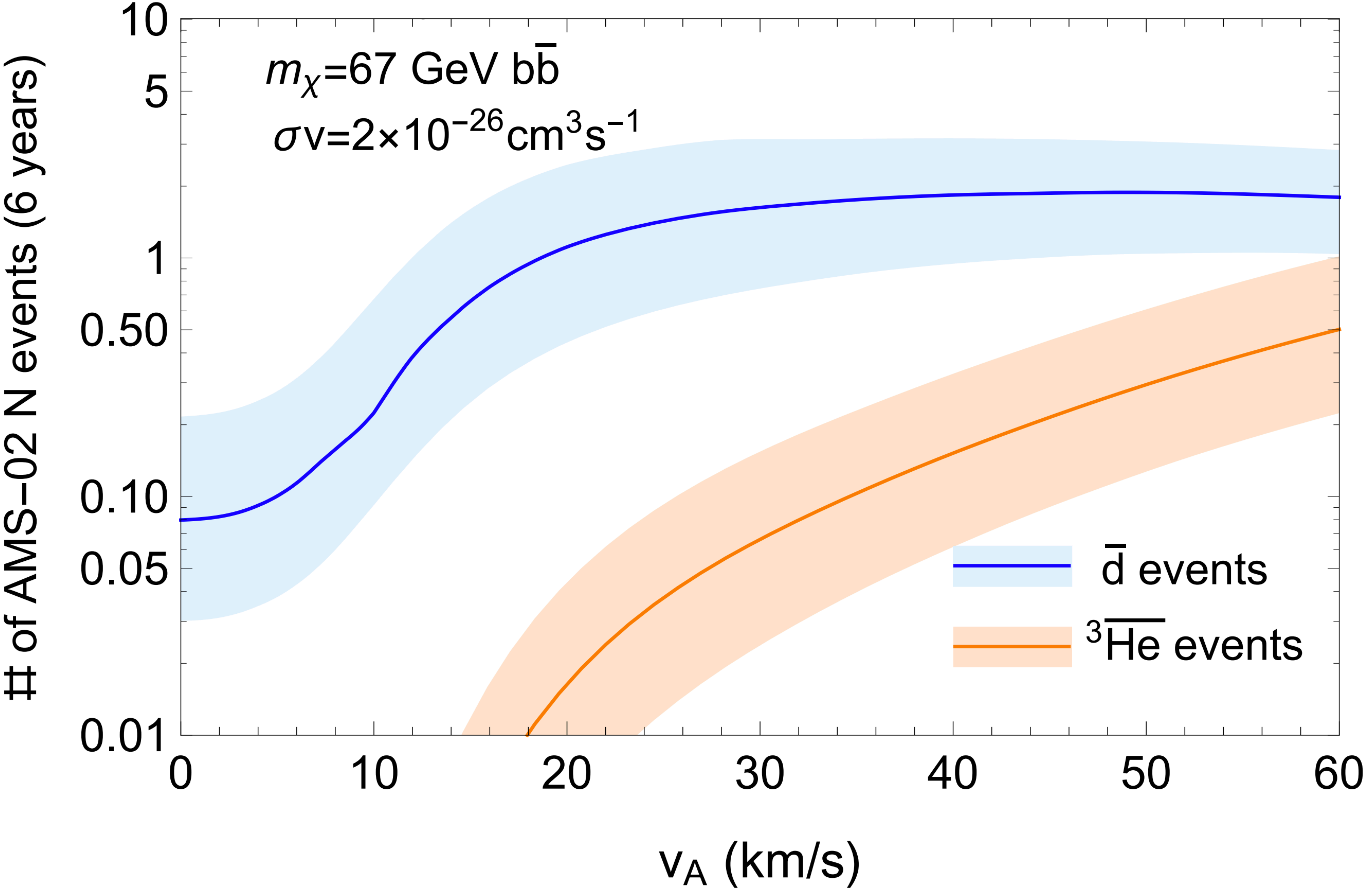
Use re-acceleration to boost the anti-He energies into the detectable range!

$$D_{pp}(R) = \frac{4}{3\delta(2-\delta)(4-\delta)(2+\delta)} \frac{R^2 v_A^2}{D_{xx}(R)}$$

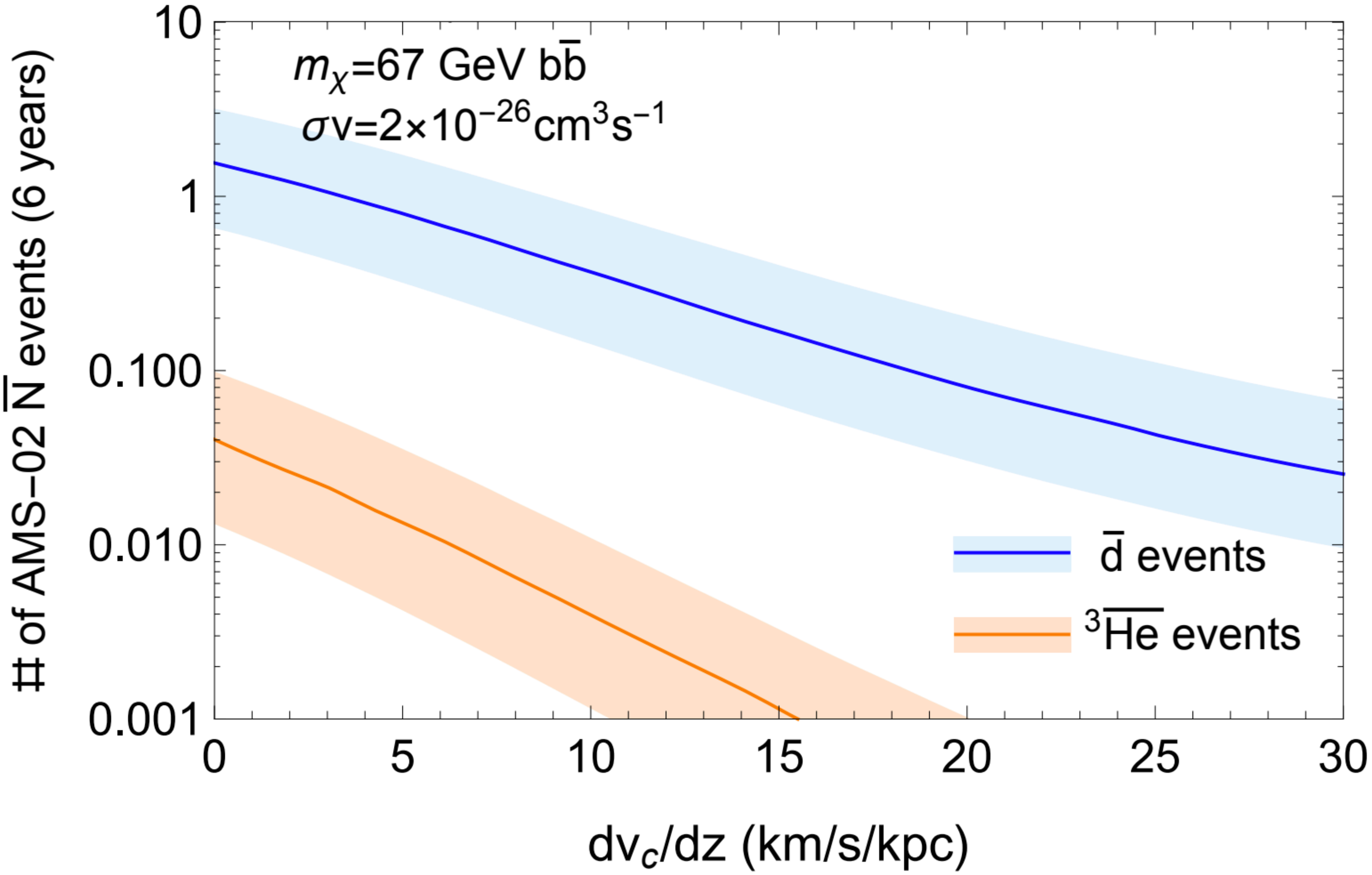




Cholis, Linden, Hooper (2020; 2001.08749)



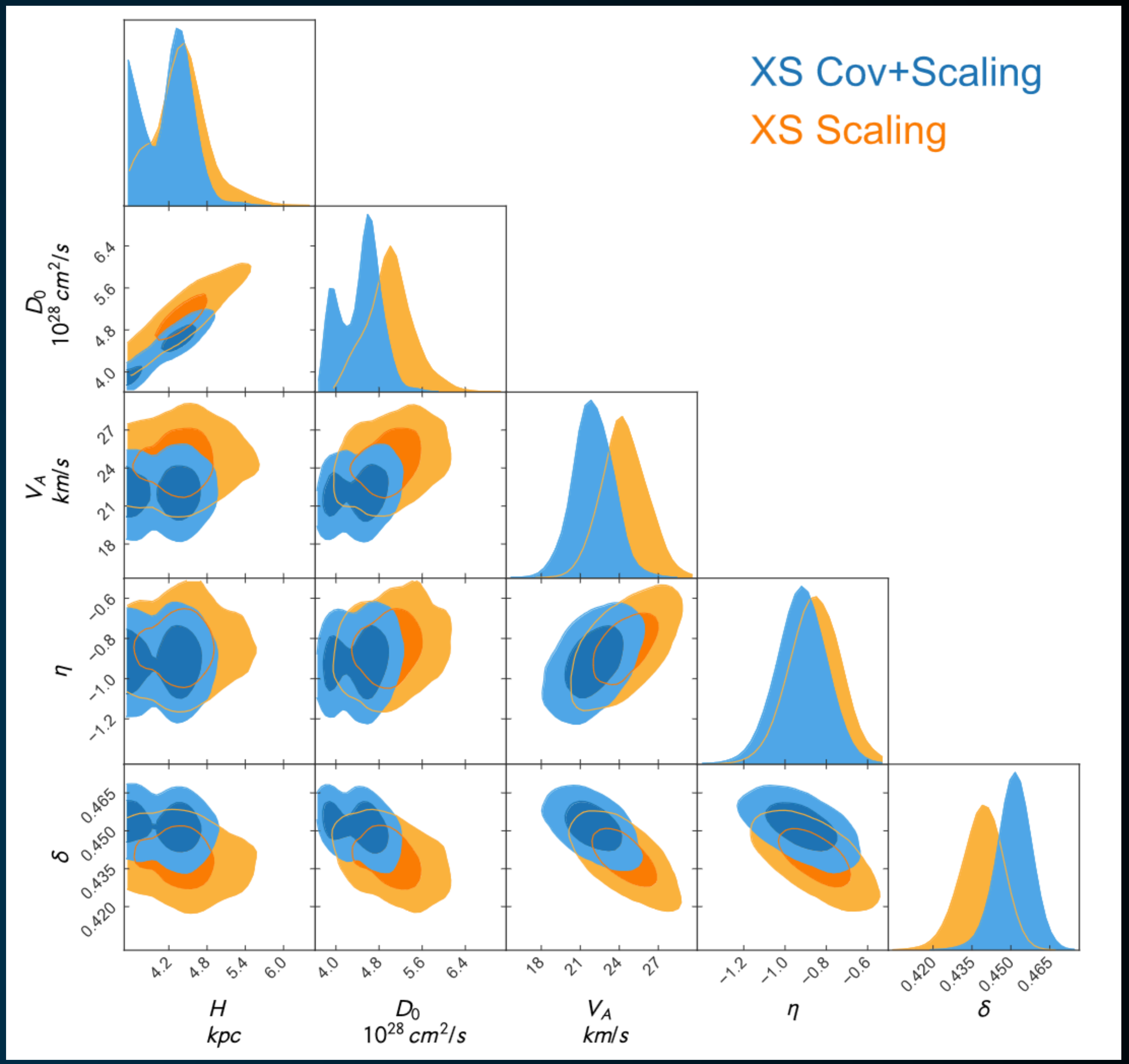
m_χ=67 GeV $b\bar{b}$
$\sigma v = 2\times10^{-26} cm^3 s^{-1}$
# of AMS−02 $\overline{N}$ events (6 years)
10
5
1
0.50
0.10
0.05
0.01
0
10
20
30
40
50
60
v_A (km/s)
$\bar{d}$ events
$^3\overline{He}$ events



10
1
0.100
0.010
0.001
0
5
10
15
20
25
30
# of AMS–02 N̄ events (6 years)
dv_c/dz (km/s/kpc)
m_χ=67 GeV bb̄
σv=2×10^−26 cm^3 s^−1
d̄ events
3He̅ events

Problem: Alfvén Velocity is Probably Not High

De la Torre Luque, TL, Winkler (2024; 2401.10329)



1.) Observations of cosmic-ray ratios strongly constrain fundamental diffusion parameters.

2.) Best-fit multi parameter models that fit e.g., B/C or heavy nuclei find v_a = 0 - 25 km/s.

- Still possibly a contributor

3.) Caveats:

- Alfvén velocity is probably not homogeneous.
- Only need high Alfvén velocity in region where anti-He is produced (e.g., GC), not in rest of spatial volume.

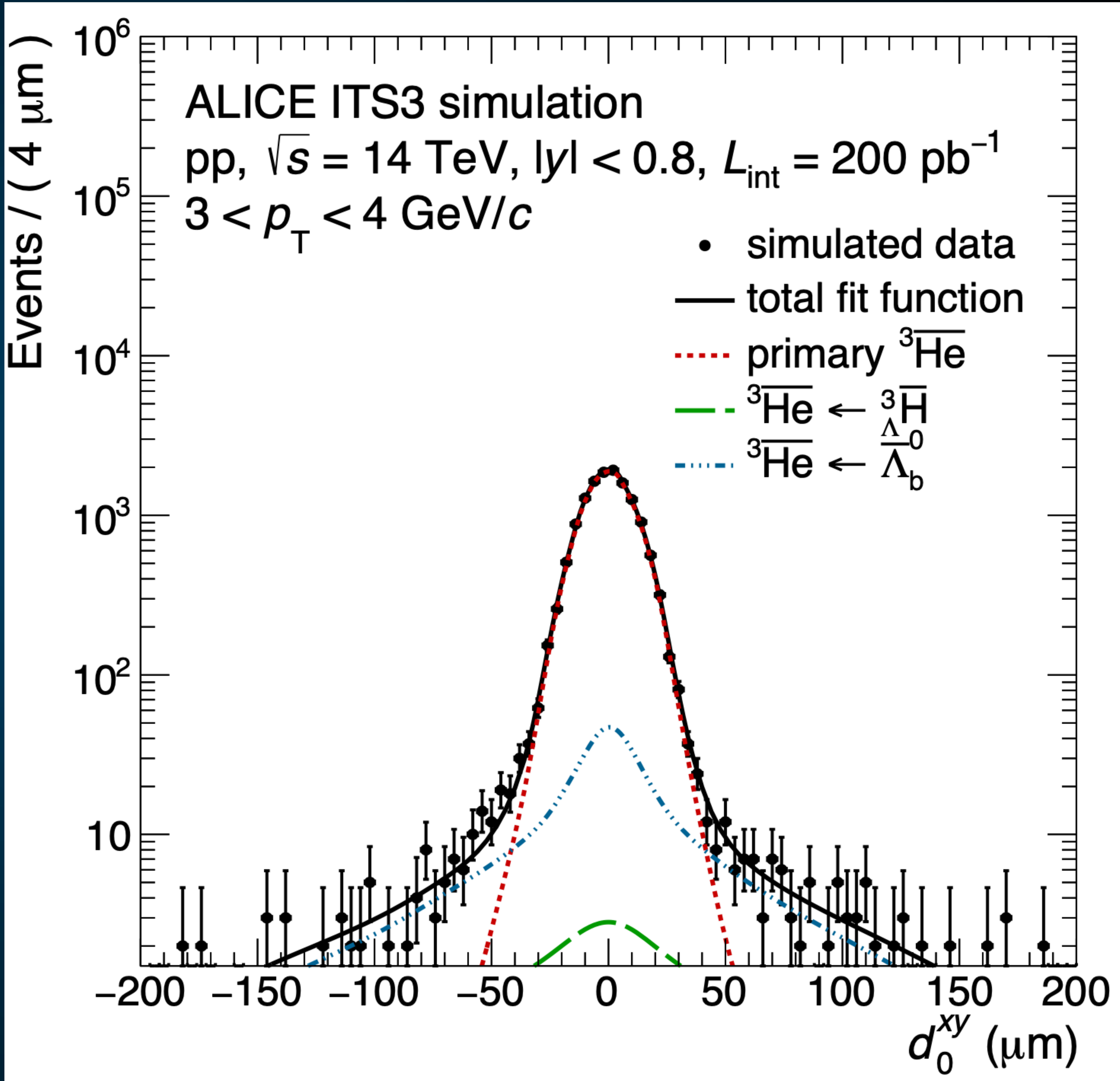
Korsmeier, Cuoco (2021; 2103.09824)

Can We Find this At Particle Accelerators?

Current observations are not sensitive to this offset

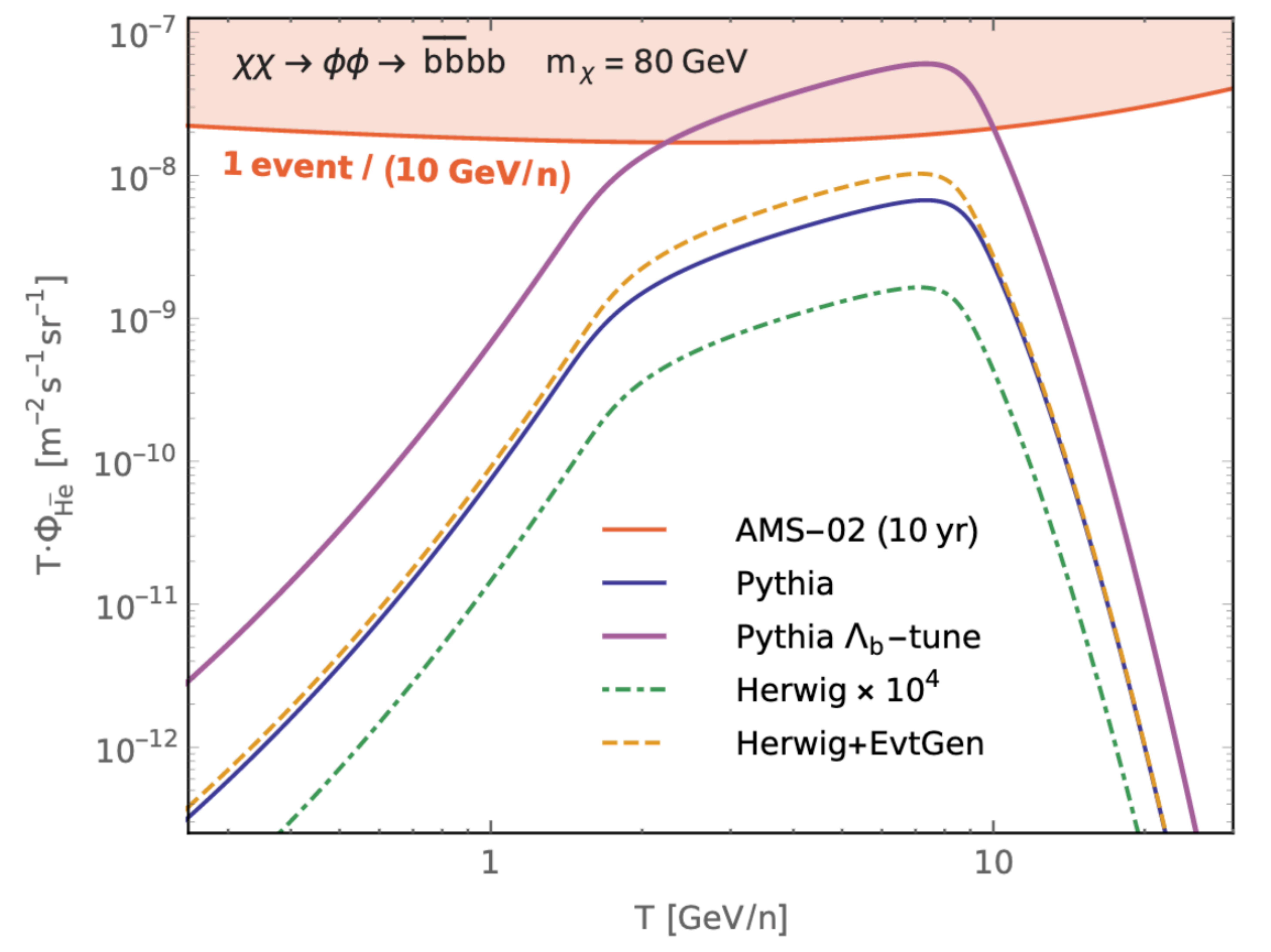
The ITS2 run of ALICE is unlikely to be able to detect the signal, but may provide a hint if the antihelium production rate is near the upper limits of our predictions.

The upcoming ITS3 experiment from ALICE will be able to differentiate the Λ_b channel



Building a Specific Dark Particle

Can further boost antihelium formation through the inclusion of a dark mediator that lies just above above the antihelium mass.



Winkler & Linden (2020; 2020.16251)

Ding, Li, & Zhou (2022; 2212.05239)

Key Insight - Coalescence Momentum for Antihelium Should Be Larger

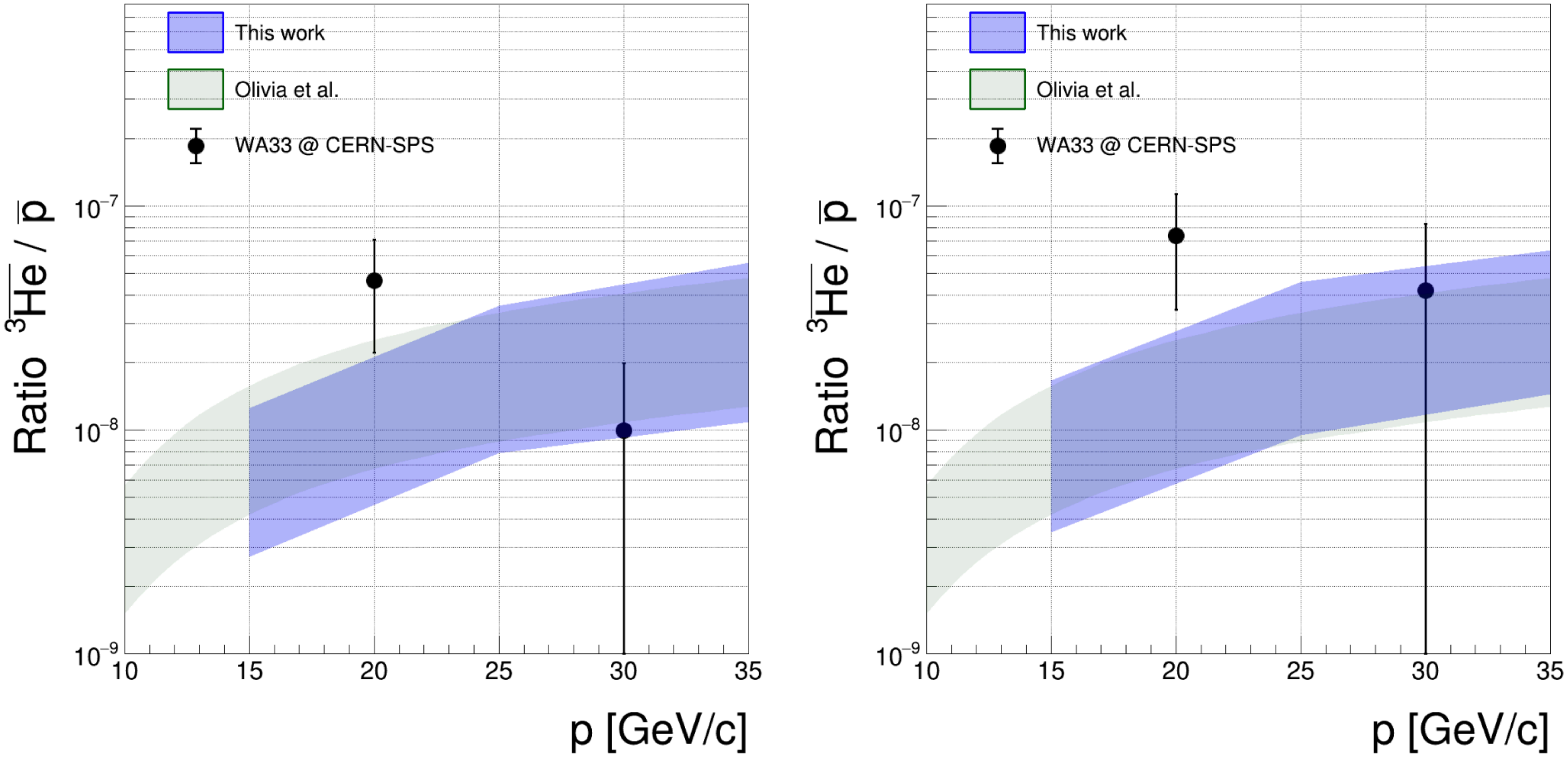


FIG. 4. The invariant production cross section ratio $^3\overline{\mathrm{He}}/\overline{p}$ as function of momentum p [GeV/c] in the laboratory frame for (left) p-Be at $p_{\mathrm{lab}} = 200\,\mathrm{GeV}/c$ and (right) p–Al at $p_{\mathrm{lab}} = 200\,\mathrm{GeV}/c$. The uncertainty bands for this work were estimated by varying the coalescence parameter from $p_{0,G}$ ($59\,\mathrm{MeV}/c$) to 130% of $p_{0,G}$ ($77\,\mathrm{MeV}/c$).

Shukla et al. (2006.12707)

Uncertainties in the Rate

1.) $\overline{\Lambda_b}$ production rate

experiment	channel	measurement	Pythia (default)	Pythia (Λ_b-tune)
LEP [4, 5]	$f(b \to \Lambda_b)$	$0.101^{+0.039}_{-0.031}$	0.037	0.101
LEP [6]	$f(b \to \Lambda_b, \Xi_b, \Omega_b)$	0.117 ± 0.021	0.047	0.127
Tevatron CDF [7]	$\frac{f(b\to\Lambda_b)}{f(b\to B)}$	$0.281^{+0.141}_{-0.103}$	0.046	0.135
LHCb [8]	$\frac{f(b\to\Lambda_b)}{f(b\to B)}$	0.259 ± 0.018	0.048	0.134

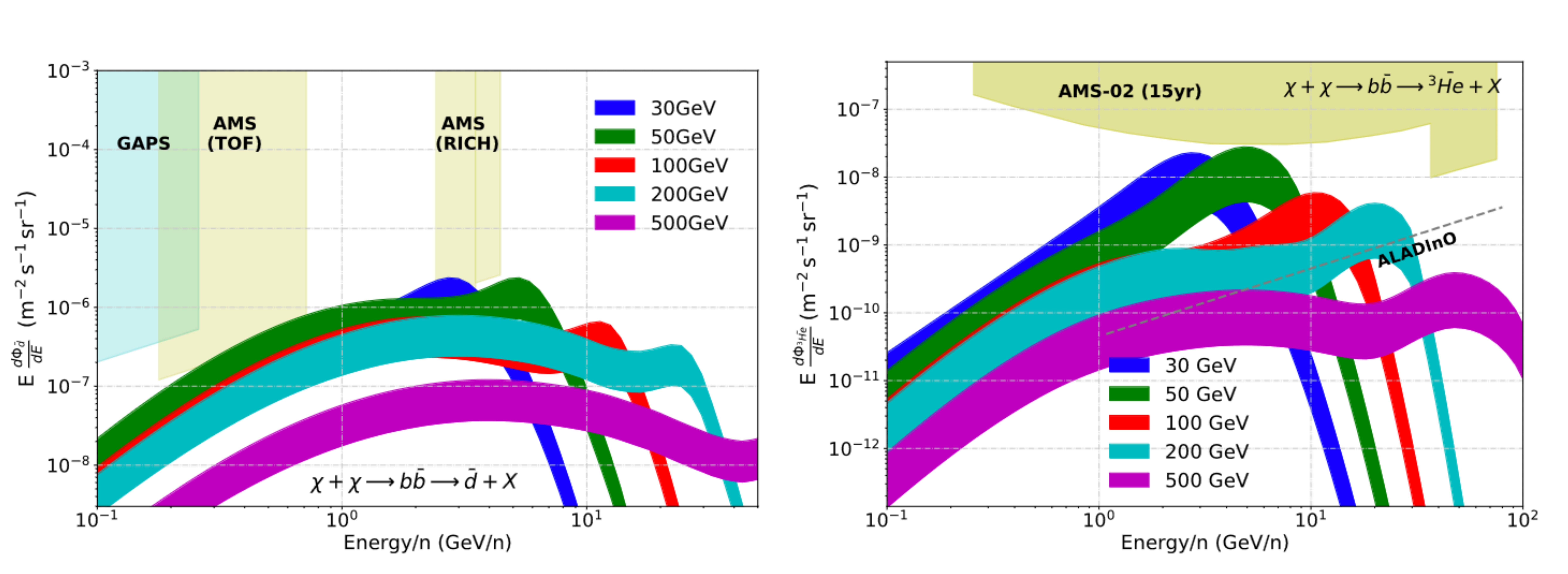
Solution: Change the parameter: $\mathrm{probQQtoQ} \to 0.24$

Enhances diquark formation and $\overline{\Lambda_b}$ production.

Winkler & Linden (2020; 2020.16251)

Strongly Coupled Dark Sectors

De la Torre Luque, Winkler, TL (2024; 2404.13114)



To Do Quantitative Predictions it is worth constraining models with antiprotons.

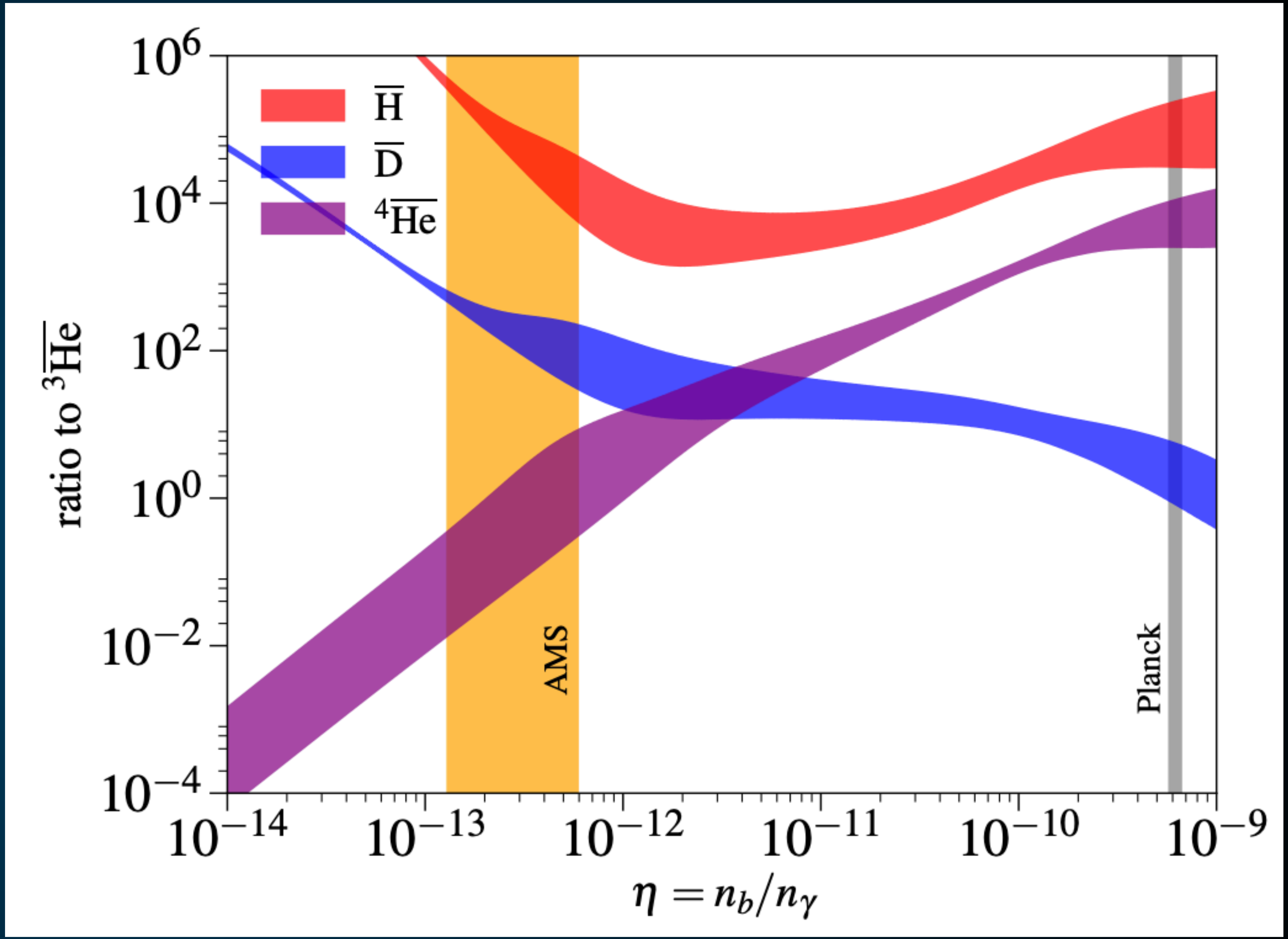
Standard models should not be seen yet, even with $\overline{\Lambda_b} \rightarrow {}^3\mathrm{He}$

Antihelium Production in Antidomains

If the big bang is asymmetric - different regions may have inverted antiparticle/ particle dominance.

Anticlouds (and potentially antistars will form), undergoing BBN and later stellar fusion.

Can produce a significant (low-energy anti helium abundance)



Poulin et al. (2018; 1808.08961)